

University of Groningen

HIFI observations of warm gas in DR21

Ossenkopf, Volker; Roellig, M.; Simon, R.; Schneider, N.; Okada, Y.; Stutzki, J.; Gerin, M.; Akyilmaz, M.; Beintema, D; Benz, A. O.

Published in:
Astronomy & Astrophysics

DOI:
[10.1051/0004-6361/201014579](https://doi.org/10.1051/0004-6361/201014579)

IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.

Document Version
Publisher's PDF, also known as Version of record

Publication date:
2010

[Link to publication in University of Groningen/UMCG research database](#)

Citation for published version (APA):

Ossenkopf, V., Roellig, M., Simon, R., Schneider, N., Okada, Y., Stutzki, J., Gerin, M., Akyilmaz, M., Beintema, D., Benz, A. O., Berne, O., Boulanger, F., Bumble, B., Coeur-Joly, O., Dedes, C., Diez-Gonzalez, M. C., France, K., Fuente, A., Gallego, J. D., ... Yorke, H. (2010). HIFI observations of warm gas in DR21: Shock versus radiative heating. *Astronomy & Astrophysics*, 518, Article L79. <https://doi.org/10.1051/0004-6361/201014579>

Copyright

Other than for strictly personal use, it is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), unless the work is under an open content license (like Creative Commons).

The publication may also be distributed here under the terms of Article 25fa of the Dutch Copyright Act, indicated by the "Taverne" license. More information can be found on the University of Groningen website: <https://www.rug.nl/library/open-access/self-archiving-pure/taverne-amendment>.

Take-down policy

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

Downloaded from the University of Groningen/UMCG research database (Pure): <http://www.rug.nl/research/portal>. For technical reasons the number of authors shown on this cover page is limited to 10 maximum.

See discussions, stats, and author profiles for this publication at: <https://www.researchgate.net/publication/45917637>

HIFI observations of warm gas in DR21: Shock versus radiative heating

Article in *Astronomy and Astrophysics* · May 2010

DOI: 10.1051/0004-6361/201014579 · Source: arXiv

CITATIONS

7

READS

25

38 authors, including:



Volker Ossenkopf
University of Cologne

228 PUBLICATIONS 3,077 CITATIONS

[SEE PROFILE](#)



M. Röllig
University of Cologne

111 PUBLICATIONS 1,004 CITATIONS

[SEE PROFILE](#)



Robert Simon
University of Cologne

192 PUBLICATIONS 2,719 CITATIONS

[SEE PROFILE](#)



Nicola Schneider
University of Cologne

279 PUBLICATIONS 5,651 CITATIONS

[SEE PROFILE](#)

Some of the authors of this publication are also working on these related projects:



Herschel-Planck [View project](#)



The Herschel Imaging survey of OB Young Stars [View project](#)

LETTER TO THE EDITOR

HIFI observations of warm gas in DR21: Shock versus radiative heating[★]

V. Ossenkopf^{1,2}, M. Röllig¹, R. Simon¹, N. Schneider³, Y. Okada¹, J. Stutzki¹, M. Gerin⁴, M. Akyilmaz¹, D. Beintema², A.O. Benz⁵, O. Berne⁶, F. Boulanger⁷, B. Bumble⁸, O. Coeur-Joly^{9,10}, C. Dedes⁵, M.C. Diez-Gonzalez¹¹, K. France¹², A. Fuente¹³, J.D. Gallego¹¹, J.R. Goicoechea¹⁴, R. Güsten¹⁵, A. Harris¹⁶, R. Higgins^{17,1}, B. Jackson², C. Jarchow¹⁸, C. Joblin^{9,10}, T. Klein¹⁵, C. Kramer¹⁹, S. Lord²⁰, P. Martin¹², J. Martin-Pintado¹⁴, B. Mookerjee²¹, D.A. Neufeld²², T. Phillips²³, J.R. Rizzo¹⁴, F.F.S. van der Tak^{2,24}, D. Teyssier²⁵, H. Yorke⁸

(Affiliations can be found after the references)

Preprint online version: May 17, 2010

ABSTRACT

Context. The molecular gas in the DR21 massive star formation region is known to be affected by the strong UV field from the central star cluster and by a fast outflow creating a bright shock. The relative contribution of both heating mechanisms is the matter of a long debate.

Aims. By better sampling the excitation ladder of various tracers we provide a quantitative distinction between the different heating mechanisms.

Methods. HIFI observations of mid- J transitions of CO and HCO⁺ isotopes allow us to bridge the gap in excitation energies between observations from the ground, characterizing the cooler gas, and existing ISO LWS spectra, constraining the properties of the hot gas. Comparing the detailed line profiles allows to identify the physical structure of the different components.

Results. In spite of the known shock-excitation of H₂ and the clearly visible strong outflow, we find that the emission of all lines up to ≥ 2 THz can be explained by purely radiative heating of the material. However, the new *Herschel*/HIFI observations reveal two types of excitation conditions. We find hot and dense clumps close to the central cluster, probably dynamically affected by the outflow, and a more widespread distribution of cooler, but nevertheless dense, molecular clumps.

Key words. ISM: structure – ISM: kinematics and dynamics – ISM: molecules – HII regions – Submillimeter

1. Introduction

DR21 is a deeply embedded H II region created by the radiation from at least six OB stars (Roelfsema et al., 1989). It sits within a ridge of dense molecular material that obscures the H II-region at optical wavelengths. The embedded cluster drives a violent bipolar outflow in north-east to south-west direction. It is prominent in the 2 μ m emission of vibrationally excited H₂, tracing hot, shocked gas and in the Spitzer 4.5 μ m channel (Garden et al., 1986; Davis et al., 2007). As the cluster is located close to the eastern edge of the molecular ridge, the eastern, blue-shifted outflow expands in a blister-like fountain, while the western, red-shifted outflow is highly collimated (Lane et al., 1990).

Spitzer 8 μ m images reveal spots of bright PAH emission with a size below 10'' (Marston et al., 2004). They represent the surfaces of high-density, UV irradiated clumps, forming photon-dominated (or photo-dissociation) regions (PDRs), transition zones from ionized and atomic gas to dense molecular gas where physics and chemistry are dominated by UV radiation from young stars (Hollenbach & Tielens, 1999; Ossenkopf et al., 2007). These hot and dense regions give rise to the emission of PDR tracers, such as HCO⁺, high- J CO, atomic and ionized carbon, and atomic oxygen.

An overview on the numerous existing observations in DR21 is given by Schneider et al. (2006) and Jakob et al. (2007). In spite of the wealth of data, the heating of the molecular gas is still

debated. The bright H₂ emission indicates shock heating in the outflow and the wings of the molecular lines prove its dynamical impact, but excitation models of the observed emission of CO, [C I], [C II], and [O I] by Lane et al. (1990) and Jakob et al. (2007) have shown that the emission of those tracers cannot be explained by shocks, but is consistent with a pure UV heating, i.e., PDR physics.

To quantify the heating of the gas, we analyze spectra taken with the HIFI instrument (de Graauw et al., 2010) on board the *Herschel Space Observatory* (Pilbratt et al., 2010) during the performance verification campaign. Observations of high- J HCO⁺ transitions trace hot material, ionized by UV radiation or X-rays (Sternberg & Dalgarno, 1995). Hot water lines are produced in shocked gas (Snell et al., 2005), and observations of CO isotopes around $J = 10$ close the gap in the excitation ladder between ground-based observations and existing ISO data, allowing to obtain a full picture of the temperature distribution.

In Sect. 2 we present the observational data. Sect. 3 compares the measured line profiles to distinguish different components based on their velocity distribution. In Sect. 4 we provide a model for the emission, supporting the PDR character of the source, and discuss the results in Sect. 5.

2. Observations

2.1. HIFI observations

All spectra presented here were obtained in performance verification observations for the HIFI instrument. As their main goal

[★] *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

Table 1. Summary of the used HIFI observational data

transition	ν_{line} [GHz]	HPBW "	observing mode ¹	$t_{\text{int,source}}$ [s]	rms [K]
HCO ⁺ 6-5	535.062	40	OTF map	16	0.04
HCO ⁺ 12-11	1069.694	21	FSW spectral scan	150	0.1
¹³ CO 10-9	1101.350	21	LC spectral scan	270	0.08
C ¹⁸ O 10-9	1097.163	21	LC spectral scan	270	0.08
[C II]	1900.537	12	DBS raster map	14	1
		20	convolved+binned	112	0.2
		40	convolved+binned	420	0.1

¹ DBS = dual-beam-switch, OTF = On-The-Fly, FSW = frequency-switch, LC = load-chop, OFF position = 20h37m10s, 42°37'00"

was to demonstrate the functionality and performance of the different observing modes, the spectra were taken with a large variety of observing modes and strategies. Consequently, every spectrum was taken in a slightly different manner. All observational parameters are summarized in Table 1.

Most observations were single-point observations towards the central position of the DR21 H II region at RA=20h39m01.1s, DEC=42°19'43.0" (J2000). Fully sampled maps were only obtained in the [C II] line. Data were taken with the wideband spectrometer (WBS) at a resolution of 1.1 MHz, corresponding to 0.2 km/s (at 1900 GHz) – 0.7 km/s (at 500 GHz). The [C II] data were rebinned to a velocity resolution of 0.45 km/s to improve the signal to noise.

2.2. Complementary data

ISO Long Wavelength Spectrometer 43 – 197 μm grating scans were obtained for the DR21 central position from the ISO Data Archive (TDT 15200786). Integrated line intensities were extracted for [O I] at 63 and 145 μm and the CO 14–13 to 17–16 transitions. Mid- J CO lines of the DR21 region were mapped with the KOSMA 3 m submm telescope (Jakob et al., 2007). We use the lines of CO and ¹³CO from $J = 3-2$ to 7-6, which have been observed at native angular resolutions from 80" to 40". The HCO⁺ 1-0, H¹³CO⁺ 1-0, and HCO⁺ 3-2 observations were taken with the IRAM 30 m telescope (Schneider et al., 2010). Native angular resolutions at the 1-0 and 3-2 transition frequencies are 28" and 9", respectively.

2.3. Beam size effects

For a direct comparison of the different data sets we smoothed the available data to the coarsest common angular resolution of 80", matching that of the CO 3-2 KOSMA beam and the lowest frequency ISO observations. This is impossible for the single-point HIFI observations. Moreover, the *Herschel* beam varies between 40" HPBW for the HCO⁺ line at 535 GHz and 20" for ¹³CO at 1100 GHz. Those data were corrected for the different beam filling by estimating the source size from the HCO⁺ 3-2 line as a PDR tracer (Sternberg & Dalgarno, 1995) with quite compact emission (20" to 30" in diameter). Successively convolving from the native angular resolution of 9" to the angular resolution of the *Herschel* data, and further to the final, smoothed spectra at 80", we obtained scaling factors for the HIFI spectra, being 0.5 when going from 40" to 80" and 0.33 when going from 20" to 80". The maps obtained in [C II] allow a direct smoothing

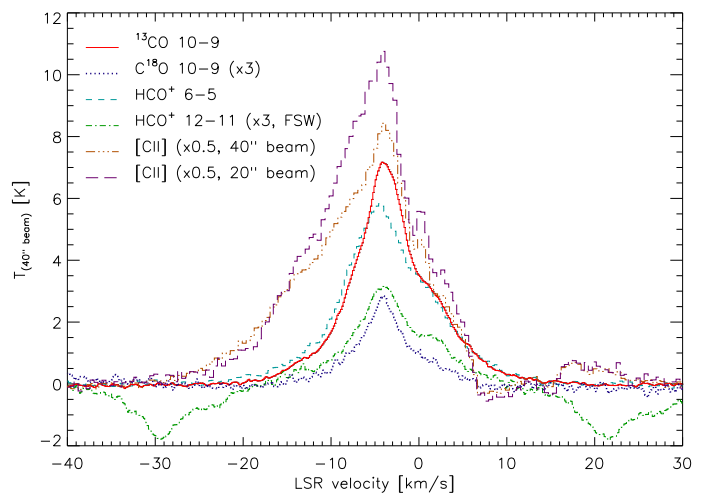


Fig. 1. HIFI spectra of the lines used in the fit of the PDR properties. The [C II] data are convolved either to the resolution of the HCO⁺ 6-5 line or to that of the CO 10-9 lines. The negative features in the HCO⁺ 12-11 line are artifacts from the frequency-switch observing mode.

to a resolution of 40" (see Table 1). Beyond that size, the same scaling factor as above was applied.

3. Line profiles

The molecular ridge including the DR21 H II region has an intrinsic LSR velocity of -3 km s^{-1} . The quiescent material is visible in narrow absorption lines of NH₃ (Matsakis et al., 1977) and H₂CO (Bieging et al., 1982). A second velocity component at 8-10 km s^{-1} is known to be associated with the W75N complex. It appears in emission in CO and [C I] (Jakob et al., 2007), as a narrow absorption feature in HCO⁺ 1-0 (Nyman, 1983), and has a very broad velocity distribution in the H I 21 cm absorption (Thompson et al., 1969; Roberts et al., 1997). The wings of the low- J CO lines trace outflow velocities down to about -20 km s^{-1} for the eastern, blister outflow and up to $\approx 10 \text{ km s}^{-1}$ for the western outflow.

Figure 1 shows the profiles of the HIFI spectra of CO isotopes, HCO⁺ and [C II]. All lines peak at about -4 km s^{-1} . The CO and HCO⁺ lines have a similar shape, but the [C II] line shows an additional broad blue wing extending down to -30 km s^{-1} . This indicates that the warm molecular material is slightly blue-shifted relative to the cold gas and that the [C II] emission is not only originating from that warm gas, but also from the ionized wind in the blister outflow.

Figure 2 compares the shapes of the CO and HCO⁺ isotope lines with complementary ground-based measurements towards the same positions. We show only a few selected transitions as, e.g., the data for the CO 7-6 or 4-3 lines provide no additional information. All CO isotopic lines up to 7-6 are roughly symmetric, centered at the ridge velocity of -3 km s^{-1} . The lines of the main isotope are heavily self-absorbed with the absorption dip marking line center. The 10-9 lines, tracing hotter material, are slightly asymmetric and shifted to -4 km s^{-1} . All HCO⁺ lines, except H¹³CO⁺ 1-0, have a profile very similar to the 10-9 lines of the CO isotopes indicating that HCO⁺ is mainly abundant in the heated layer and less in the overall molecular material. The H¹³CO⁺ 1-0 line has a deviating profile centered at -2 km s^{-1} tracing global infall (Kirby, 2009). CO 3-2 and HCO⁺ 1-0 show an additional feature at the W75N complex velocity.

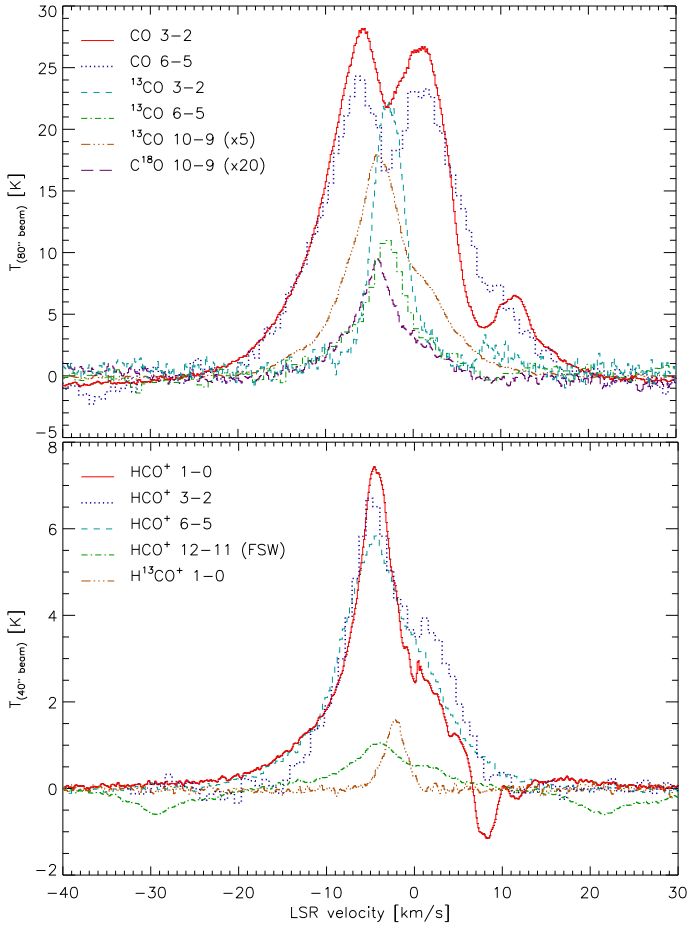


Fig. 2. Selected profiles of CO isotope lines (upper plot) and HCO⁺ isotope lines towards the DR21 central position. All lines are convolved or scaled to the coarsest common resolution (see Sect.2.3).

To better understand the exact velocity distribution in the source, we plot some additional HIFI lines towards the same position (see Falgarone et al., 2010; Ossenkopf et al., 2010; van der Tak et al., 2010) in Fig. 3. The 1097 GHz line of hot water matches the rest of the PDR tracers. The foreground at -3 km s^{-1} is visible in absorption in the ground state transitions of water, but in emission in a nearby OFF position for [C II]. At velocities of about $+2 \text{ km s}^{-1}$, we find the warm gas facing the redshifted western outflow, apparent as a secondary peak in the ground state water and CH⁺ lines and as a shoulder in ¹³CO and the hot water line. Finally, we can clearly identify the cold W75N component at $\approx -9 \text{ km s}^{-1}$.

For the warm gas we can distinguish two velocity components – a blue shifted one related to the blister outflow at -4 km s^{-1} and a second one, at 2 km s^{-1} , related to the western outflow. This is consistent with the clumpy PDR geometry proposed by Lane et al. (1990). In terms of line modelling, it is, however, impossible to separate the two outflow directions as their emission is overlapping to a large degree. We simply add their intensities. The [C II] emission requires a special treatment because of the additional emission from the ionized outflow. To take this into account, we have obtained an integrated [C II] intensity by scaling the ¹³CO 10-9 line profile to match the peak and the red wing of the [C II] profile and ignoring the remaining blue-wing emission.

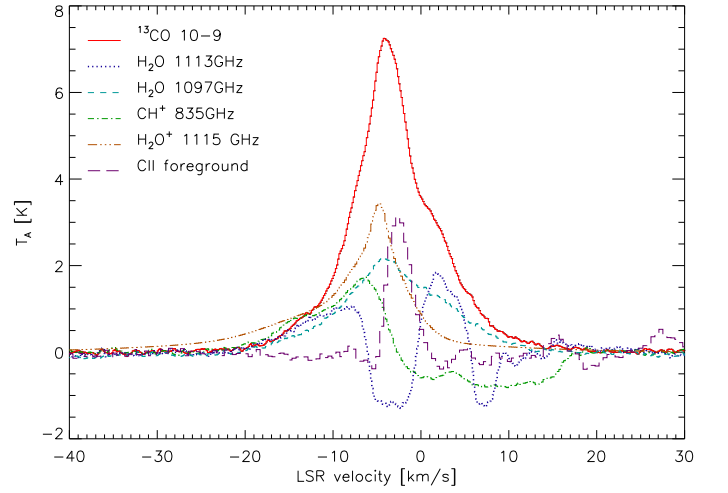


Fig. 3. Auxiliary lines measured by HIFI towards the same position. They can be used to estimate the contribution of the foreground material. The ¹³CO 10-9 profile is also displayed for comparison.

4. Modelling

We use the KOSMA- τ PDR code (Röllig et al., 2006) to model the emission of PDR ensembles, representing a distribution of spherical clumps with $dN/dM \propto M^{-1.8}$ (Cubick et al., 2008). For DR21 two ensembles with different properties had to be superimposed, a hot component, close to the inner H II region with strong FUV illumination, but only a small fraction of the total mass, and a cooler component that fills a larger solid angle and provides the bulk of the material. Each clumpy PDR ensemble has five free parameters: the average ensemble density, n_{ens} , the ensemble mass, M_{ens} , the UV field strength, χ given in units of the Draine field, and the minimum and maximum mass of the clump ensemble, $[M_{\text{min}}, M_{\text{max}}]$. In contrast to most other PDR models, we fit absolute line intensities, using the available ground-based observations, complementary ISO data, and the HIFI lines of the CO isotopes, HCO⁺, atomic and ionized carbon, and atomic oxygen. The chemical network that has been applied in these calculations includes ¹³C but not ¹⁸O. The C¹⁸O lines were scaled from the ¹³CO intensities with a conversion of 1:8. Simulated annealing was used to find the optimum parameter combination.

The significance of the model is limited by the fact that the clump superposition ignores mutual line shading between different clumps, i.e., optical depth effects are only considered within individual clumps. This is usually justified by the virialised velocity dispersion between different clumps, but for optically very thick and broad lines some correction is needed. To estimate the effect we have computed the optical depth for the bulk of the individual clumps. This is of the order of unity or below for the majority of the observed transitions, reaches values up to ten for ¹³CO and HCO⁺ transitions up to $J = 4$, but exceeds ten for the CO main isotope lines up to $J = 6$ and the [O I] 63 μm line. For the CO lines showing clear self-absorption dips, we performed a Gaussian fit to the line wings and used the integrated intensity of that Gaussian to compute the total emission including the blocked radiation from the inner clumps close to the H II-region. As we have no spectral information for the [O I] line, we have no estimate for the blocked radiation in this case, so that we excluded that data point from the fit.

The best fit result is shown in Figure 4. The corresponding model parameters are:

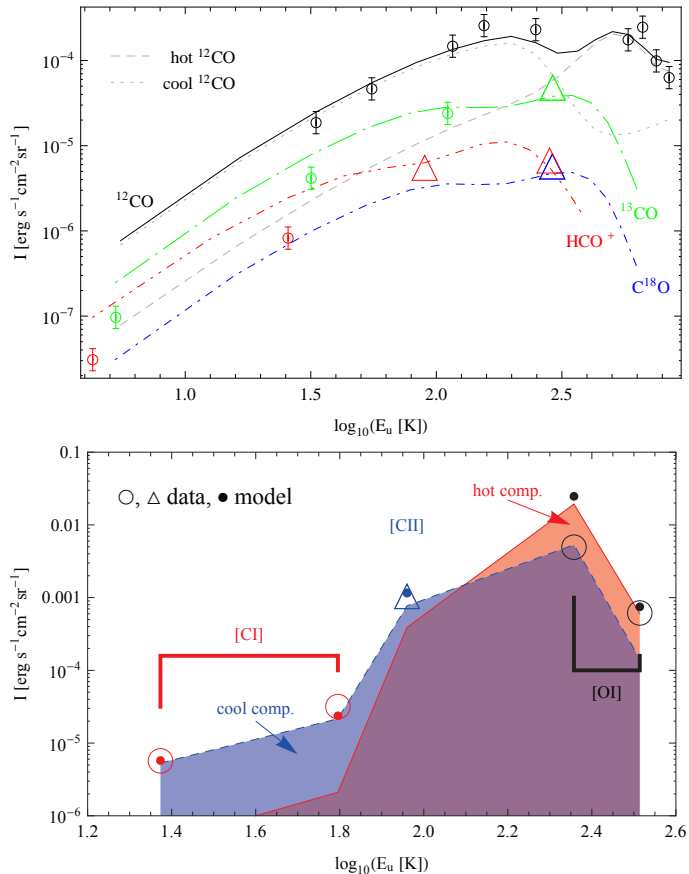


Fig. 4. Two-ensemble PDR model fit to the observed CO, HCO⁺, and fine structure line intensities, shown as function of the upper level energy. HIFI measurements are depicted as open triangles, complementary data points as open circles. The dotted line and the shaded areas indicate the contributions from the two ensembles to the CO and the fine structure lines, respectively.

Ensemble 1: $\chi = 1.0 \cdot 10^5$, $n_{\text{ens}} = 1.3 \cdot 10^6 \text{cm}^{-3}$, $M_{\text{ens}} = 150 M_{\odot}$,
 $[M_{\text{min}}, M_{\text{max}}] = [10^{-2}, 8 \cdot 10^1] M_{\odot}$

Ensemble 2: $\chi = 3.2 \cdot 10^2$, $n_{\text{ens}} = 1.1 \cdot 10^6 \text{cm}^{-3}$, $M_{\text{ens}} = 830 M_{\odot}$,
 $[M_{\text{min}}, M_{\text{max}}] = [10^{-3}, 10^1] M_{\odot}$

The two-ensemble model provides a reasonable fit to all 22 data points. The model slightly overestimates the emission for the optically thick rotational lines of ¹³CO and HCO⁺ which showed no clear self-absorption dip, so that we have insufficient optical depth corrections. The reduced χ^2 value of the fit amounts to 11.1, where a contribution of 9.5 results from the three optically thick low- J data points of ¹³CO and HCO⁺. For the remaining 19 data points we would obtain an excellent reduced χ^2 value of 1.6. The topology of the χ^2 function shows several side minima, but they are worse by at least a factor two. As expected, the model predicts a too high [O I] intensity, as it ignores that the outer clumps of the cooler ensemble block the contribution from the hot inner component.

5. Discussion

The double-peak structure of the predicted CO intensities across the excitation ladder reflects the two different UV fields leading to different excitation conditions. No single-parameter ensemble can fit all data, but distinct low temperature and high temperature components are needed. The parameters of the model are in agreement with independent estimates. A UV flux of 10^5 Draine

fields corresponds to a geometrical distance of 0.06 pc, i.e. 7'', from the central cluster, matching the size of the PAH emission (Marston et al., 2004). Our clump densities match those determined by Jakob et al. (2007) for the extended cool gas, but are slightly higher than their hot-gas density ($4 \cdot 10^5 \text{cm}^{-3}$). In contrast, Jones et al. (1994); van der Tak et al. (2010) find still somewhat higher densities for the hot gas, up to 10^7cm^{-3} . The hot ensemble mass is close to the $170 M_{\odot}$ derived from early CO 7-6 observations by Jaffe et al. (1989). The total mass of the PDR ensembles falls between the mass limits derived by Jakob et al. (2007) from dust observations and from line radiative transfer fits.

While the existing ground-based observations provide a very good constraint on the properties of the extended cool gas, and the ISO lines show the total amount of hot gas, it is only the set of new HIFI data that puts the hot and cold distributions well apart from each other in terms of the temperature structure. While Jakob et al. (2007) obtained cooling curves with single peaks, the new data for the 10-9 lines of the CO isotopes and the HCO⁺ transitions force the fit to a bimodal distribution of excitation conditions. When we exclude the *Herschel* data from the model fit, we obtain a parameter set that shows a UV field that is lower by a factor ten for the hot ensemble, i.e., that would imply molecular clumps farther away from the central cluster. For the cold ensemble, the fitted UV field is also somewhat lower, while all other parameters remain similar to those from the full fit. Only with the *Herschel* data, we therefore obtain a parameter set that is consistent with the source geometry.

As the two-ensemble PDR model is able to fit all of the observed lines, we find no evidence for a shock heating of the dense gas. This is in agreement with the analysis of Lane et al. (1990), explicitly excluding a shock origin of the fine-structure lines, but seems to be in contradiction with the analysis of the line profiles in Sect. 3 that shows excited outflow material. We conclude that the material visible in the blue line wing, characterizing the blister outflow, is contained in dense clumps that are accelerated by the outflow, but that are chemically and energetically fully dominated by the UV field and not by the associated shock.

References

- Bieging, J.H.; Wilson, T.L.; Downes, D. 1982, A&ASS 49, 607
 Cubick, M.; Stutzki, J.; Ossenkopf, V.; Kramer, C.; Röllig, M. 2008, A&A 488, 623
 Davis, C. J., Kumar, M. S. N., Sandell, G., et al. 2007, MNRAS 374, 29
 Falgarone, E.; Ossenkopf, V.; Gerin, M. et al. 2010, A&A this volume
 Garden, R., Geballe, T. R., Gatley, I., Nadeau, D. 1986, MNRAS, 220, 203
 de Graauw, Th.; et al. 2010, this volume
 Jaffe, D.T., Genzel, R., Harris, A.I., Lugten, B., Stacey, G.J., Stutzki, J., ApJ 344, 265
 Hollenbach, D.J.; Tielens, A.G.G.M. 1999, Rev. Mod. Phys. 71, 173
 Jakob, H.; Kramer, C.; Simon, R.; Schneider, N.; Ossenkopf, V.; Bontemps, S.; Graf, U.U.; Stutzki, J. 2007, A&A 461, 999
 Jones, K. N.; Field, D.; Gray, M. D.; Walker, R. N. F. 1994, A&A 288, 581
 Kirby, L. 2009, ApJ 694, 1056
 Lane, A.P.; Haas, M.R.; Hollenbach, D.J.; Erickson, E.F. 1990, ApJ 361, 132
 Marston, A. P., Reach, W. T., Noriega-Crespo, A., et al. 2004, ApJS, 154, 333
 Matsakis, D. N.; Brandshaft, D.; Chui, M.F. et al. 1977, ApJ 214, L67
 Ossenkopf, V.; Röllig, M.; Cubick, M.; Stutzki, J. 2007, in *Molecules in Space and Laboratory*, by J.L. Lemaire & F. Combes (eds.), S. Diana publ., p.95
 Ossenkopf, V.; Müller, H.S.P.; Lis, D. 2010, A&A this volume
 Nyman, L.-Å. 1983, A&A 120, 307
 Pilbratt, G.; et al. 2010, this volume
 Roberts, D.A.; Dickel, H.R.; Goss, W.M. 1997, ApJ 476, 209
 Roelfsema, P. R., Goss, W. M., Geballe, T. R. 1989, A&A, 222, 247
 Röllig, M.; Ossenkopf, V.; Jeyakumar, S.; Stutzki, J.; Sternberg, A. 2006, A&A 451, 917
 Schneider, N.; Bontemps, S.; Simon, R.; et al. 2006, A&A 458, 855

- Schneider, N., Csengeri, T., Bontemps, S. 2010, A&A accepted, arXiv:1003.4198
- Snell R.L., Hollenbach D., Howe J.E., Neufeld D.A., Kaufman M.J., Melnick G.J., Bergin E.A., Wang Z. 2005, ApJ 620, 758
- Sternberg, A., Dalgarno, A. 1995, ApJS, 99, 565
- Thompson, A.R.; Colvin, R.S., Hughes, M.P. 1969, ApJ 158, 939
- van der Tak, F. F. S.; Marseille, M. G.; Herpin, F. 2010, A&A this volume
-
- ¹ I. Physikalisches Institut der Universität zu Köln, Zülpicher Straße 77, 50937 Köln, Germany
- ² SRON Netherlands Institute for Space Research, P.O. Box 800, 9700 AV Groningen, Netherlands
- ³ Laboratoire AIM, CEA/DSM - INSU/CNRS - Université Paris Diderot, IRFU/Sap CEA-Saclay, 91191 Gif-sur-Yvette, France
- ⁴ LERMA & UMR 8112 du CNRS, Observatoire de Paris and École Normale Supérieure, 24 Rue Lhomond, 75231 Paris Cedex 05, France
- ⁵ Institute for Astronomy, ETH Zürich, 8093 Zürich, Switzerland
- ⁶ Leiden Observatory, Universiteit Leiden, P.O. Box 9513, NL-2300 RA Leiden, The Netherlands
- ⁷ Institut d'Astrophysique Spatiale, Université Paris-Sud, Bât. 121, 91405 Orsay Cedex, France
- ⁸ Jet Propulsion Laboratory, 4800 Oak Grove Drive, MC 302-231, Pasadena, CA 91109 U.S.A.
- ⁹ Université de Toulouse, UPS, CESR, 9 avenue du colonel Roche, 31028 Toulouse cedex 4, France
- ¹⁰ CNRS, UMR 5187, 31028 Toulouse, France
- ¹¹ Observatorio Astronómico Nacional (IGN), Centro Astronómico de Yebes, Apartado 148. 19080 Guadalajara
- ¹² Department of Astronomy and Astrophysics, University of Toronto, 60 St. George Street, Toronto, ON M5S 3H8, Canada
- ¹³ Observatorio Astronómico Nacional (OAN), Apdo. 112, 28803 Alcalá de Henares (Madrid), Spain
- ¹⁴ Centro de Astrobiología, CSIC-INTA, 28850, Madrid, Spain
- ¹⁵ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69, 53121, Bonn, Germany
- ¹⁶ Astronomy Department, University of Maryland, College Park, MD 20742, USA
- ¹⁷ Experimental Physics Dept., National University of Ireland: Maynooth, Co. Kildare, Ireland
- ¹⁸ MPI für Sonnensystemforschung, D37191 Katlenburg-Lindau, Germany
- ¹⁹ Instituto de Radio Astronomía Milimétrica (IRAM), Avenida Divina Pastora 7, Local 20, 18012 Granada, Spain
- ²⁰ IPAC/Caltech, MS 100-22, Pasadena, CA 91125, USA
- ²¹ Tata Institute of Fundamental Research (TIFR), Homi Bhabha Road, Mumbai 400005, India
- ²² Department of Physics and Astronomy, Johns Hopkins University, 3400 North Charles Street, Baltimore, MD 21218, USA
- ²³ California Institute of Technology, 320-47, Pasadena, CA 91125-4700, USA
- ²⁴ Kapteyn Astronomical Institute, University of Groningen, P.O. Box 800, 9700 AV Groningen, Netherlands
- ²⁵ European Space Astronomy Centre, Urb. Villafraanca del Castillo, P.O. Box 50727, Madrid 28080, Spain

Acknowledgements. HIFI has been designed and built by a consortium of institutes and university departments from across Europe, Canada and the United States under the leadership of SRON Netherlands Institute for Space Research, Groningen, The Netherlands and with major contributions from Germany, France and the US. Consortium members are: Canada: CSA, U.Waterloo; France: CESR, LAB, LERMA, IRAM; Germany: KOSMA, MPIfR, MPS; Ireland: NUI Maynooth; Italy: ASI, IFSI-INAF, Osservatorio Astrofisico di Arcetri- INAF; Netherlands: SRON, TUD; Poland: CAMK, CBK; Spain: Observatorio Astronómico Nacional (IGN), Centro de Astrobiología (CSIC-INTA). Sweden: Chalmers University of Technology - MC2, RSS & GARD; Onsala Space Observatory; Swedish National Space Board, Stockholm University - Stockholm Observatory; Switzerland: ETH Zurich, FHNW; USA: Caltech, JPL, NHSC.

This work was supported by the German *Deutsche Forschungsgemeinschaft*, DFG project number Os 177/1-1. A portion of this research was performed at the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.