

University of Groningen

Time lags of the kilohertz quasi-periodic oscillations in the low-mass X-ray binaries 4U 1608-52 and 4U 1636-53

de Avellar, Marcio G. B.; Mendez, Mariano; Sanna, Andrea; Horvath, Jorge E.

Published in:
Monthly Notices of the Royal Astronomical Society

DOI:
[10.1093/mnras/stt1001](https://doi.org/10.1093/mnras/stt1001)

IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.

Document Version
Publisher's PDF, also known as Version of record

Publication date:
2013

[Link to publication in University of Groningen/UMCG research database](#)

Citation for published version (APA):

de Avellar, M. G. B., Mendez, M., Sanna, A., & Horvath, J. E. (2013). Time lags of the kilohertz quasi-periodic oscillations in the low-mass X-ray binaries 4U 1608-52 and 4U 1636-53. *Monthly Notices of the Royal Astronomical Society*, 433(4), 3453-3463. <https://doi.org/10.1093/mnras/stt1001>

Copyright

Other than for strictly personal use, it is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), unless the work is under an open content license (like Creative Commons).

The publication may also be distributed here under the terms of Article 25fa of the Dutch Copyright Act, indicated by the "Taverne" license. More information can be found on the University of Groningen website: <https://www.rug.nl/library/open-access/self-archiving-pure/taverne-amendment>.

Take-down policy

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

Downloaded from the University of Groningen/UMCG research database (Pure): <http://www.rug.nl/research/portal>. For technical reasons the number of authors shown on this cover page is limited to 10 maximum.

Time lags of the kilohertz quasi-periodic oscillations in the low-mass X-ray binaries 4U 1608–52 and 4U 1636–53

Marcio G. B. de Avellar,¹* Mariano Méndez,² Andrea Sanna² and Jorge E. Horvath¹

¹*Instituto de Astronomia, Geofísica e de Ciências Atmosféricas, Universidade de São Paulo, Rua do Matão 1226, 05508-090 São Paulo, Brazil*

²*Kapteyn Astronomical Institute, University of Groningen, PO Box 800, 9700 AV Groningen, the Netherlands*

Accepted 2013 June 4. Received 2013 May 28; in original form 2013 February 26

ABSTRACT

We studied the energy and frequency dependence of the Fourier time lags and intrinsic coherence of the kilohertz quasi-periodic oscillations (kHz QPOs) in the neutron-star low-mass X-ray binaries 4U 1608–52 and 4U 1636–53, using a large data set obtained with the *Rossi X-ray Timing Explorer*. We confirmed that, in both sources, the time lags of the lower kHz QPO are soft and their magnitude increases with energy. We also found that: (i) In 4U 1636–53, the soft lags of the lower kHz QPO remain constant at $\sim 30 \mu\text{s}$ in the QPO frequency range 500–850 Hz, and decrease to $\sim 10 \mu\text{s}$ when the QPO frequency increases further. In 4U 1608–52, the soft lags of the lower kHz QPO remain constant at $40 \mu\text{s}$ up to 800 Hz, the highest frequency reached by this QPO in our data. (ii) In both sources, the time lags of the upper kHz QPO are hard, independent of energy or frequency and inconsistent with the soft lags of the lower kHz QPO. (iii) In both sources the intrinsic coherence of the lower kHz QPO remains constant at ~ 0.6 between 5 and 12 keV, and drops to zero above that energy. The intrinsic coherence of the upper kHz QPO is consistent with being zero across the full energy range. (iv) In 4U 1636–53, the intrinsic coherence of the lower kHz QPO increases from ~ 0 at ~ 600 Hz to ~ 1 , and it decreases to ~ 0.5 at 920 Hz; in 4U 1608–52, the intrinsic coherence is consistent with the same trend. (v) In both sources the intrinsic coherence of the upper kHz QPO is consistent with zero over the full frequency range of the QPO, except in 4U 1636–53 between 700 and 900 Hz where the intrinsic coherence marginally increases. We discuss our results in the context of scenarios in which the soft lags are either due to reflection off the accretion disc or up-/down-scattering in a hot medium close to the neutron star. We finally explore the connection between, on one hand the time lags and the intrinsic coherence of the kHz QPOs, and on the other the QPOs' amplitude and quality factor in these two sources.

Key words: stars: neutron – X-rays: binaries – X-rays: individual: 4U 1608–52 – X-rays: individual: 4U 1636–53.

1 INTRODUCTION

The kilohertz quasi-periodic oscillations (kHz QPOs) are the fastest variability so far observed from any astrophysical object. These QPOs are relatively narrow peaks in the power density spectra of neutron-star (NS) low-mass X-ray binaries (LMXBs). These peaks often appear in pairs at frequencies ν_1 and $\nu_2 > \nu_1$ that change with time, ranging from ~ 400 up to ~ 1300 Hz (see van der Klis 2006, and references therein; we will refer to the kHz QPO at ν_1 and ν_2 as the lower and the upper kHz QPO, respectively). Since the Keplerian frequency close to the innermost stable circular orbit (ISCO) around an ~ 15 km and $\sim 1.4 M_\odot$ NS is ~ 1300 Hz, several

models have been proposed in which the kHz QPOs, and especially the one at ν_2 , reflect the Keplerian frequency of matter orbiting at the inner edge of an accretion disc, very close to the NS surface (e.g. Miller, Lamb & Psaltis 1998; Stella & Vietri 1999).

Most of the observational and theoretical studies in the past 15 years concentrated on the frequencies (e.g. Psaltis, Belloni & van der Klis 1999; van Straaten et al. 2002; Altamirano et al. 2008; Sanna et al. 2012) and, to a lesser degree, the amplitude and the coherence of the kHz QPOs (Méndez, van der Klis & Ford 2001; Di Salvo, Méndez & van der Klis 2003; Barret et al. 2005a; Barret, Olive & Miller 2005b, 2006; Méndez 2006; Sanna et al. 2010, see van der Klis 2006 for a review). Almost no work has been done on another aspect of the kHz QPO signal: the energy-dependent time (or phase) lags and the Fourier coherence (Nowak et al. 1999) of the kHz QPOs.

*E-mail: marcavel@astro.iag.usp.br

The time (phase) lags and Fourier coherence are Fourier-frequency-dependent measures of, respectively, the time (phase) delay and the degree of linear correlation between two concurrent and correlated signals, in this case light curves of the same source, in two different energy bands. Perfect (unit) intrinsic (noise-subtracted) coherence entails that one can predict one of the signals from the measurements of the other (see Nowak et al. 1999 for a detailed explanation of the time lags, coherence and intrinsic coherence, and how to calculate them). There are only two papers in the literature that describe observational results of the time lags of the kHz QPOs: Using data from the *Rossi X-ray Timing Explorer (RXTE)*, Vaughan et al. (1997, see also Vaughan et al. 1998 for an erratum) studied the time lags of the kHz QPOs in the NS LMXBs 4U 1608–52, 4U 1636–53 and 4U 0614+09, while Kaaret et al. (1999) did a similar work for the NS LMXB 4U 1636–53. These authors found that in all these sources, within a narrow frequency interval centred around the centroid frequency of the QPOs, the low-energy (soft) photons lag the high-energy (hard) photons by ~ 5 – $100 \mu\text{s}$, with the magnitude of the lags increasing as the energy difference between soft and hard photons increases. Those two works were done at the beginning of the *RXTE* mission, when only a handful of observations for each of these sources were available; furthermore, the analysis had to be done using a coarse energy resolution to alleviate the limited signal-to-noise ratio of the data, and in some cases the lags could not be measured accurately or significantly. More importantly, at the time it was not possible to study the potential dependence of the time lags upon the frequency of the kHz QPO. Surprisingly, no similar observational study has been done since then, despite the fact that the *RXTE* archive contains enough data that could be used to explore this.

Here, we extend the initial work by Vaughan et al. (1997, 1998) and Kaaret et al. (1999), using a large set of *RXTE* observations to analyse the energy-dependent time lags and coherence of the kHz QPOs in two of the three NS LMXBs mentioned above: 4U 1608–52 and 4U 1636–53. 4U 1608–52 is a transient NS LMXB showing outbursts with recurrence times that vary from 80 d to a few years (Lochner & Roussel-Dupre 1994). kHz QPOs have been detected in this source with frequencies ranging from ~ 500 to ~ 1050 Hz (Berger et al. 1996; Méndez et al. 1998b; van Straaten, van der Klis & Méndez 2003). Vaughan et al. (1997, 1998) measured time lags in one of the kHz QPO (later on identified as the lower kHz QPO; see Méndez et al. 1998b) in this source, going from $\sim -10 \mu\text{s}$ at ~ 7 keV to $\sim -60 \mu\text{s}$ at ~ 23 keV. (All time lags were calculated with respect to photons in the 4–6 keV energy band. Because of this convention, a negative lag means that the hard photons lead the soft ones. These are usually called soft lags, as opposed to hard lags when the hard photons lag the soft ones.) 4U 1636–53 is a persistent NS LMXB that shows kHz QPOs with very similar properties (frequency range, amplitude and coherence) to the kHz QPOs in 4U 1608–52 (Méndez, van der Klis & Paradijs 1998c; Di Salvo et al. 2003; Barret et al. 2006). Kaaret et al. (1999) measured soft lags in the lower kHz QPO in this source ranging from $\sim -5 \mu\text{s}$ at ~ 7 keV to $\sim -40 \mu\text{s}$ at ~ 23 keV, relative to the 3.8–6.4 keV band. The only measurement of a time lag of the upper kHz QPO is that by Vaughan et al. (1997) who found a time lag of $-26 \pm 23 \mu\text{s}$ between the 6.5–67 keV and 2–6.5 keV photons for the upper kHz QPO in 4U 0614+091.

The origin of the time lags in the power density spectra of black hole (Miyamoto et al. 1988; Nowak et al. 1999) and NS (Ford et al. 1999) LMXBs, and supermassive black holes in active galactic nuclei (Zoghbi et al. 2010) is not fully understood. The mechanisms proposed so far involve Compton up-/down-scattering of photons

produced in the accretion disc (and in the case of NS, on the NS surface or boundary layer) in a corona of hot electrons that surrounds the system (Payne 1980; Sunyaev & Titarchuk 1980; Lee & Miller 1998; Lee, Misra & Taam 2001; Falanga & Titarchuk 2007) or, in the case of soft lags in active galactic nuclei, reflection of hard photons from the corona off the inner parts of the accretion disc (Zoghbi et al. 2010; Zoghbi, Uttley & Fabian 2011). If these mechanisms are applicable to NS systems, the time lags of the kHz QPOs provide a powerful tool to study the physics of the accretion flow in the vicinity of the NS, since these lags encode information about the size and the geometry of the scattering medium or the reflector. From the implied light-crossing time, the observed time lags must be produced within a very small region close to the NS surface. Here, we combine *RXTE* data of each of these two sources separately, using a technique (see Section 2) that allows us to increase the signal-to-noise ratio of the time lags and Fourier coherence, and at the same time enables us to study, for the first time, the dependence of the time lags upon the frequency of the kHz QPO. We present our results in Section 3, and discuss them in the context of existing models in Section 4.

2 OBSERVATIONS AND DATA ANALYSIS

For this paper, we used data of 4U 1608–52 and 4U 1636–53 taken with the Proportional Counter Array (PCA; Jahoda et al. 2006) onboard *RXTE* (Bradt, Rothschild & Swank 1993). For each observation (see below), we calculated the complex Fourier transform in a number of energy bands. In all cases, we calculated fast Fourier transforms up to a Nyquist frequency of 2048 Hz over segments of 16 s, yielding a minimum frequency and a frequency resolution of $1/16$ Hz. For each of these 16-s segments, we stored the real and imaginary part of the Fourier transform as a function of frequency. We did this for all available observations.

For the analysis of 4U 1608–52, we used data from the 1998 outburst of the source. These are the same observations used by Méndez et al. (1998b). From those data, we only used the observations in which the source showed kHz QPOs (observations marked with circles in fig. 1 of Méndez et al. 1998b). The data were taken using two different observing modes in which the full energy band of the PCA (nominally from 2 to 60 keV) was either divided into 32 or 64 energy channels (32M and 64M; see Table 1). Due to the ageing of the PCA instrument, the channel-to-energy conversion of the PCA changed over time. Those changes, however, took place over relatively long time-scales (of the order of years; see for instance van Straaten et al. 2003), except for a few occasions in which the

Table 1. Mean energy of the different energy bands used to calculate the time lags of the kHz QPOs in 4U 1608–52 and for 4U 1636–53. In the case of 4U 1608–52, we show separately the mean energy in the two observing modes (32M and 64M; see Section 2).

	4U 1608–52		4U 1636–53
\bar{E} (keV)	(32M)	\bar{E} (keV) (64M)	\bar{E}
	3.75	3.75	4.2
	5.56	5.42	6.0
	7.96	7.34	8.0
	10.94	9.47	10.2
	14.33	12.29	12.7
	17.49	15.86	16.3
	–	18.10	18.9

mission control changed the gain of the PCA instrument manually.¹ Because the observations of 4U 1608–52 that we used here span a time interval of less than 30 d, and during that period no manual gain changes were applied to the PCA, we could use always the same channel scheme to define the energy bands in our analysis (Table 1).

In the case of 4U 1636–53, we used all the data available in the *RXTE* archive for this source up to 2011 May; these are the same data used by Sanna et al. (2012). Here, we only selected the 528 observations of 4U 1636–53 in which Sanna et al. (2012) found at least one of the kHz QPOs in the power density spectrum of this source. The *RXTE* archive contains observations of 4U 1636–53 in different modes that sampled the full PCA band in different ways and with different number of channels. These observations were carried out over a period of about 15 years, during which the gain of the instrument changed significantly (both due to manual changes as well as to the ageing of the instrument itself). We therefore had to adjust our channel selections for different groups of observations, depending on the epoch of the observation, to always select more or less the same energy bands for our analysis. We give the average energy in each of the bands that we used for our analysis of both sources in Table 1.

We used the results of Méndez et al. (1998b) and Sanna et al. (2012), who measured the frequency of the lower kHz QPO over different time segments, to assign a lower QPO frequency to each 16-s segment for which we calculated the Fourier transform. When those authors measured the lower kHz QPO frequency over a segment longer than 16 s, we assigned the same lower QPO frequency to all our 16-s segments that were included within their time segment. We further assigned the same lower QPO frequency to the same 16-s time segment for all the Fourier transforms calculated in the different energy bands.

We also assigned an upper QPO frequency to each 16-s segment. We did this as follows: For the case of 4U 1608–52 we used the plot in fig. 3 of Méndez et al. (1998b) that shows $\Delta\nu = \nu_2 - \nu_1$ versus ν_1 for the observations that we used in this paper. For each point in that figure, we added $\nu_1 + \Delta\nu$ to obtain ν_2 for the corresponding value of ν_1 . We then assigned that upper kHz QPO frequency to all 16-s segments that were used by Méndez et al. (1998b) to get those measurements of ν_1 and $\Delta\nu$. For the case of 4U 1636–53, we used the measurements of Sanna et al. (2012) of the upper kHz QPO per individual observation (typically 3000–5000 s exposure time), and we assigned that upper QPO frequency to all 16-s segments that belonged to that observation. As in the case of the lower kHz QPO, for each source we assigned the same upper QPO frequency to the same 16-s time segment in all the Fourier transforms calculated in the different energy bands.

In summary, we assigned a lower and an upper kHz QPO frequency to each 16-s segment for which we calculated the Fourier transform, the same lower and upper kHz QPO frequency per 16-s segment for each of the energy bands shown in Table 1.

In our observations of 4U 1608–52, the lower and upper kHz QPO covered the frequency range of ~ 540 – 820 Hz and ~ 860 – 1060 Hz, respectively. In the case of 4U 1636–53, the lower and the upper kHz QPO covered the frequency range of ~ 550 – 980 Hz and ~ 440 – 1250 Hz, respectively. We therefore defined a number of frequency intervals for the lower and the upper kHz QPO for each source; we show the selected frequency intervals in Table 2. We

chose the frequency intervals such that they spanned a sufficiently small frequency range to minimize the effect of possible frequency-dependent changes of the time lags of the kHz QPO, but large enough to have a sufficiently large number (balanced among the frequency intervals) of 16-s segments in each frequency interval to reduce the statistical uncertainties.

For each source, and for each kHz QPO separately, we applied the shift-and-add technique (Méndez et al. 1998b) to check the significance of the QPO in the power density spectra for different selections of the data. In both sources, both kHz QPOs were highly significant when we produced a single power density spectrum from all observations with kHz QPOs combined, shifting and adding the power density spectra of the individual time segments such that the frequency of either the lower or the upper kHz QPO was always the same. Similarly, both kHz QPOs were highly significant in both sources when we created shifted-and-added power density spectra for each frequency interval defined in Table 2, combining the photons of all energy bands, as well as when we created shifted-and-added power density spectra for each energy band defined in Table 1 combining the data of all the frequency intervals in Table 2. We did not always detect the kHz QPOs significantly when we created power density spectra dividing the data both in energy and frequency. Given that we have ~ 20 times more data for 4U 1636–53 than for 4U 1608–52, this problem occurred more often in the latter than in the former source. In general, we detected the kHz QPOs more significantly when we selected energies between ~ 5 and ~ 16 keV and the frequency of the lower and the upper kHz QPO were, respectively, in the range ~ 600 – 850 Hz and ~ 500 – 1000 Hz (see Berger et al. 1996; Barret et al. 2006). For those frequency and energy selections for which we did not detect either of the kHz QPOs significantly, we measured the time lags and intrinsic coherence of the QPOs (see below) within a frequency range centred around the QPO that we detected in other cuts of the data. By doing this, we implicitly assumed that the QPO frequency is independent of energy (see e.g. Berger et al. 1996).

We then applied the same idea behind the shift-and-add technique introduced by Méndez et al. (1998b) to study the kHz QPOs in Fourier power density spectra, but this time to study the complex Fourier spectra, and in particular, the time lags and the intrinsic coherence of the kHz QPOs in these two sources. To calculate the time lags and the intrinsic coherence of the kHz QPOs, we followed the same procedure described by Vaughan et al. (1997) and Kaaret et al. (1999), but in this case we first aligned the frequency axes of the Fourier transforms of each 16-s segment before we calculated the average cross spectra. To do this, for each frequency interval and for each source separately, we first shifted the frequency scale of the Fourier spectra such that in each frequency selection first the frequency of the lower and then that of the upper kHz QPO was always the same. For each source, we also calculated the average frequency of the QPO for each frequency interval (Table 2), separately for the lower and the upper kHz QPO. We then selected one energy band from the ones given in Table 1 as the reference band, and we calculated average cross-spectra between each of the other energy bands in Table 1 and this reference band. It is custom to choose the lowest energy band available as the reference band, but since the time lag is a relative quantity, one is free to choose any band for that. The accuracy with which one can measure the time lags depends upon the intensity and the strength of the variability of the signal (Vaughan & Nowak 1997); therefore, it is convenient to choose a reference band in which both the variability and the source intensity are high. While it is easier to choose always the same reference band, in our case we were limited by the fact

¹ See the channel-to-energy conversion table for the PCA at http://heasarc.gsfc.nasa.gov/docs/xte/r-c_table.html.

Table 2. Frequency intervals, mean frequency, and number of 16-s segments for the lower and the upper kHz QPO in 4U 1608–52 and 4U 1636–53. In the case of 4U 1608–52, we show separately the number of 16-s segments in the two observing modes (32M and 64M).

4U 1608–52			4U 1636–53			
ν range (Hz)	$\bar{\nu}$ (Hz) (32M 64M)	Lower kHz QPO		ν range (Hz)	$\bar{\nu}$ (Hz)	Number of 16-s segments
		Number of 16-s segments (32M 64M)				
540–610	579 –	635 –		550–620	596	7615
610–690	655 677	553 238		620–670	650	9131
690–770	745 735	582 399		670–715	693	9315
770–823	791 783	616 309		715–750	731	7263
–	–	–		750–790	768	7462
–	–	–		790–820	807	8802
–	–	–		820–850	836	10 011
–	–	–		850–880	864	11 907
–	–	–		880–910	895	16 076
–	–	–		910–975	921	12 100
Upper kHz QPO						
859– 910	887 –	611 –		440– 540	492	7052
910– 993	952 977	549 176		560– 650	605	4120
993–1040	1035 1028	604 353		650– 750	702	5840
1040–1064	1047 1044	622 417		750– 810	781	10 871
–	–	–		810– 870	840	8368
–	–	–		870– 930	897	8544
–	–	–		930–1025	979	11 685
–	–	–		1070–1165	1117	7577
–	–	–		1165–1250	1211	12 595

that the two sources in this paper were observed using different observing modes, and the instrument calibration changed with time (see above); this prevented us from defining exactly the same energy bands for the analysis of both sources. We therefore chose the reference band as follows, unless otherwise indicated: For the lower kHz QPO in 4U 1608–52 all time lags are calculated relative to the band centred at 3.75 keV, both for data observed in the 32M and 64M mode. For the upper kHz QPO in 4U 1608–52, we used the band centred at 7.96 keV for the 32M data and the band centred at 7.34 keV for the 64M data. To simplify the description, we will refer to the bands just mentioned as the band centred at ~ 7.5 keV. For the lower and the upper kHz QPO in 4U 1636–53, we used the bands centred at 4.2 and 10.2 keV, respectively.

Finally, to calculate the time lags and intrinsic coherence of the lower kHz QPO, we averaged the energy- and frequency-dependent time lags and intrinsic coherence over a frequency interval of 16 Hz centred around the centroid frequency of the lower kHz QPO; this range corresponds to about twice the full width at half-maximum (FWHM) of the lower kHz QPO in the combined shifted-and-added power density spectrum of all observations with a lower kHz QPO. Similarly, to calculate the time lags and intrinsic coherence of the upper kHz QPOs, we averaged the energy- and frequency-dependent time lags and intrinsic coherence over a frequency interval of 140 Hz centred around the centroid frequency of the upper kHz QPO; this range corresponds to ~ 1.2 times the FWHM of the upper kHz QPO in the combined shifted-and-added power density spectrum of all observations with an upper kHz QPO.² Formally, these are the time lags and intrinsic coherence of the light curves in the two energy

bands used in the calculation, over the time-scale (the inverse of the Fourier frequency) of either the lower or the upper kHz QPO. For simplicity, in the rest of the paper, we call these the time lags and intrinsic coherence of either the lower or the upper kHz QPO.

3 RESULTS

In the upper panel of Fig. 1, we plot the time lags as a function of Fourier frequency in 4U 1608–52 for the photons around 7.96 keV with respect to the photons around 3.75 keV, for the lower kHz QPO in the frequency interval between 690 and 770 Hz. The plot only shows the frequency range around 700 Hz, and it reveals that the time lags have large statistical fluctuations (see the large error bars; throughout the paper we always quote 1σ errors), except in the interval between ~ 695 and ~ 705 Hz, around the frequency that we used as reference to shift the lower kHz QPOs. The region where the time lags have small errors spans about twice the FWHM of the lower kHz QPO in the shifted-and-added power density spectrum of this selection. Note that the time lags around the frequency of the QPO are negative, following the convention mentioned in Section 1 that negative values indicate soft lags.

In the lower panel of Fig. 1, we plot the intrinsic Fourier coherence as a function of Fourier frequency for 4U 1608–52, calculated using the same two energy bands and for the same frequency interval of the lower kHz QPO, 690 and 770 Hz, as in the upper panel of this figure. As in the case of the upper panel of this figure, it is apparent that the intrinsic coherence is not well defined (it shows large fluctuations and very large error bars), except in a narrow range around the frequency that we used as reference to shift the lower kHz QPO. The coherence is well defined (it has small error bars) at around a value of 1 within about twice the FWHM of the kHz QPO.

In Fig. 2, we plot the time lags as a function of energy for the lower (left-hand panel) and the upper kHz QPO (right-hand panel), both for 4U 1608–52 (upper panels) and 4U 1636–53 (lower panels).

² Under the assumption that the time lags and intrinsic coherence in that frequency range follow a Gaussian distribution, we calculated the error-weighted average of the time lags and the intrinsic coherence over the indicated frequency range. This is justified given the large number of segments used in each selection (see Table 2).

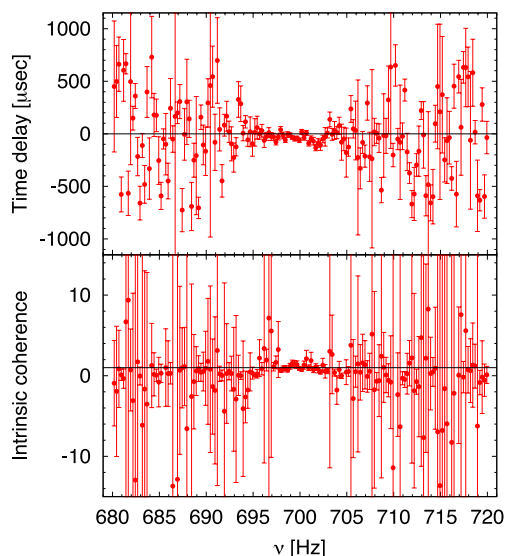


Figure 1. Upper panel: time lags as a function of Fourier frequency in 4U 1608–52, showing only the frequency range around 700 Hz. We calculated the time lags for photons with energies around 7.96 keV with respect to photons with energies around 3.75 keV (see Table 1 for the definition of the energy bands), when the frequency of the lower kHz QPO was between 690 and 770 Hz (see Table 2 for the definition of the frequency intervals). To calculate the time lags, we first shifted the frequency scale of the 16-s Fourier cross-spectra in the selected frequency interval such that the lower kHz QPO was always centred at 700 Hz, and then calculated the average cross-spectrum of all 16-s segments that belonged to that selection. The horizontal line is at a zero time lag. Lower panel: intrinsic Fourier-coherence as a function of Fourier frequency in 4U 1608–52, calculated for the same energy bands and using the same selection of frequencies of the lower kHz QPO as in the upper panel. As in the upper panel, for this plot we also applied the shift-and-add technique to align the lower kHz QPO to a reference frequency of 700 Hz. The horizontal line shows the perfect intrinsic coherence of 1.

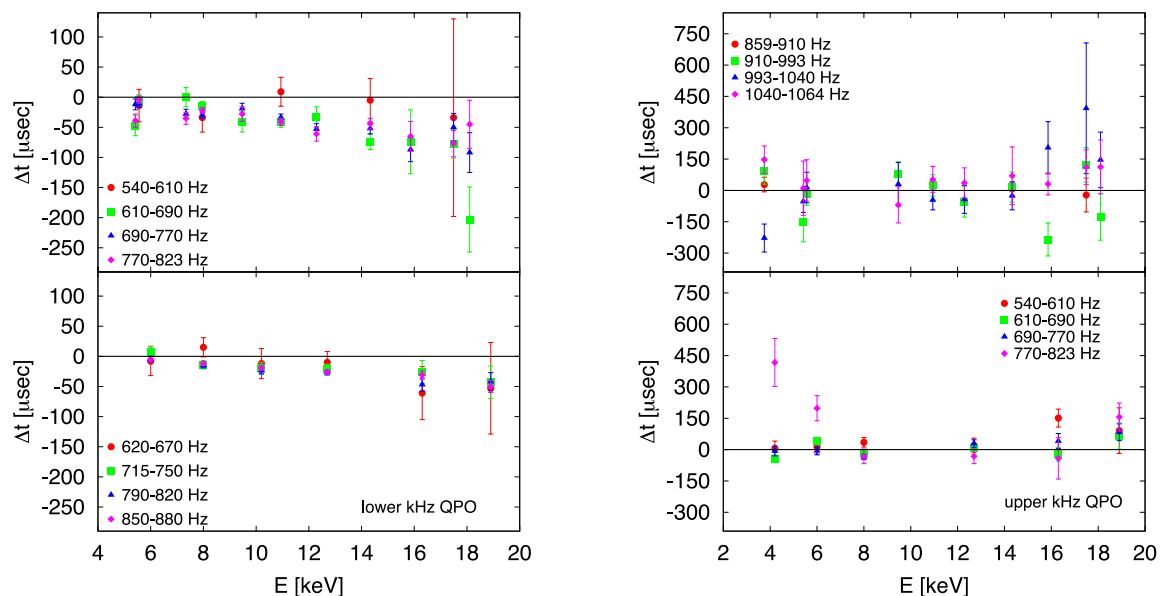


Figure 2. Time lags versus energy of the lower (left-hand column) and the upper (right-hand column) kHz QPO in 4U 1608–52 (upper row) and 4U 1636–53 (lower row). We calculated the time lags between photons with energies given by the value on the x -axis relative to photons with energies around 3.75 and ~ 7.5 keV for the lower and the upper kHz QPO in 4U 1608–52, respectively, and 4.2 and 10.2 keV for the lower and the upper kHz QPO in 4U 1636–53, respectively. The different colours and symbols show the time lags when the frequency of the QPO was within the intervals indicated in the figures. The horizontal line in all panels is at a zero time lag.

The points with different colours and symbols correspond to the time lags as a function of energy when the lower or the upper kHz QPO covered the frequency intervals indicated in the legend of the figure. In all cases, we calculated the time lags of the photons with average energy indicated in the x -axis with respect to photons in the reference band given in the previous section. For the lower kHz QPO in 4U 1608–52 (upper-left panel), the soft photons lag the hard photons (which, according to the convention we use in this paper, yields negative lags in the plot) by ~ 10 – 50 μ s at ~ 6 keV, up to ~ 100 – 200 μ s at ~ 18 keV (see Vaughan et al. 1997, 1998). For the lower kHz QPO in 4U 1636–53 (lower-left panel), the soft photons lag the hard ones by ~ 10 μ s at ~ 6 keV up to ~ 50 μ s at ~ 19 keV (see Kaaret et al. 1999). In both sources, the time lags of the upper kHz QPO (right-hand panels) have large errors, and are consistent both with zero lag and with the same trend with energy seen for the lower kHz QPOs. From this figure, it is also apparent that in both sources, and for each QPO separately, at each energy the time lags are consistent with being more or less the same in all frequency intervals, although the errors are relative large.

In the two panels of Fig. 3, we plot the time lags of the lower and the upper kHz QPO in 4U 1608–52 and 4U 1636–53 for photons with energies 12–20 keV with respect to the photons with energies 4–12 keV as a function of the average QPO frequency in each frequency interval. The upper panel shows that the time lags of the lower kHz QPO are soft (the 12–20 keV photons lead the 4–12 keV photons) in both sources. In the case of 4U 1636–53, the time lags of the lower kHz QPO are consistent with being constant at ~ -25 μ s from ~ 500 up to ~ 850 Hz, and then the magnitude of the lags decreases slightly, but significantly, to ~ -10 μ s as the frequency of the QPO increases further. The magnitude of the time lag of the last point, at ~ 920 Hz, appears to increase again but, given the error of that measurement, the value is also consistent with the overall trend of the previous points. A fit with a constant to the time lags over the full frequency range yields an average time delay of -21.0 ± 0.6 μ s, but the fit is not good, yielding $\chi^2 = 39.5$ for

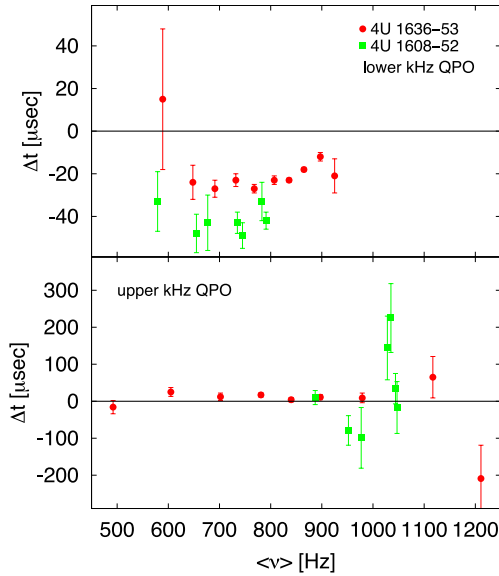


Figure 3. Time lags versus QPO frequency for photons with energy in the 12–20 keV band with respect to photons with energy in the 4–12 keV band for the lower (upper panel) and the upper kHz QPO (lower panel) of 4U 1608–52 (green squares) and 4U 1636–53 (red circles). The horizontal line in both panels is at a zero time lag.

9 degrees of freedom (dof). The F-test probability for a fit with a linear function compared to a fit with a constant is 0.012. There is no significant improvement either if we fitted with a parabola which, compared to a linear function, also yields an F-test probability of 0.012. A fit to the data with a parabola is marginally better than a fit with a constant, F-test probability of 0.002. The time lags of the lower kHz QPO in 4U 1608–52 are consistent with being constant, with an average time delay of $-42.9 \pm 0.3 \mu\text{s}$ ($\chi^2 = 3.1$ for 6 dof), albeit with larger error bars than in the case of 4U 1636–53, but since the measurements do not extend beyond ~ 800 Hz, we cannot discard a similar trend as the one observed in 4U 1636–53 at the high end of the frequency range.

The lower panel of Fig. 3 shows that in 4U 1636–53 the time lags of the upper kHz QPO are positive (the 12–20 keV photons lag the 4–12 keV photons) and independent of frequency. The average time lag of the upper kHz QPO in 4U 1636–53 over the whole frequency range is $11 \pm 3 \mu\text{s}$ ($\chi^2 = 12.9$ for 8 dof), almost 4σ different from zero. The average time lag of the upper kHz QPO in 4U 1608–52 over the full frequency range is $5 \pm 15 \mu\text{s}$ ($\chi^2 = 14.9$ for 6 dof), both consistent with the average time lag in 4U 1636–53 (over a narrower frequency range, 890–1040 Hz), and with zero lags. However, in both sources the average time lag of the upper kHz QPO is significantly different than the average time lag of the lower kHz QPO, and in 4U 1636–53, for which the statistics are better, it is apparent that the time lags of the lower and the upper kHz QPOs have opposite signs. In other words, while for the lower kHz QPO the soft photons lag the hard ones, for the upper kHz QPO the hard photons lag the soft ones.

Since the time lags of the lower and the upper kHz QPO, both in 4U 1608–52 and 4U 1636–53, appear not to change much with QPO frequency (except for the increase of the lags of the lower kHz QPO in 4U 1636–53 mentioned above; see Figs 2 and 3), for each of the kHz QPOs and for each source, we calculated the energy-dependent time lags over the whole frequency range spanned by each kHz QPO (i.e. no frequency selection). The upper panel of

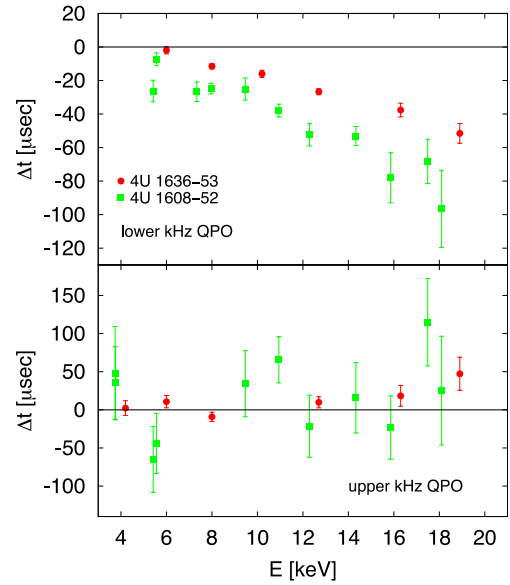


Figure 4. Time lags versus energy for the lower and the upper kHz QPOs in 4U 1608–52 (green squares) and 4U 1636–53 (red circles), for all frequency intervals combined. Upper panel: time lags of the lower kHz QPO for photons with energies given by the value on the x-axis with respect to photons with energies around 3.75 and 4.2 keV for 4U 1608–52 and 4U 1636–53, respectively. Lower panel: time lags of the upper kHz QPO for photons with energies given by the value on the x-axis with respect to photons with energies around ~ 7.5 and 10.2 keV for 4U 1608–52 and 4U 1636–53, respectively. The horizontal line in both panels is at a zero time lag.

Fig. 4 shows that the magnitude of the soft time lags of the lower kHz QPO increases more or less linearly with energy in both sources (the best-fitting linear relation yields $\chi^2 = 9.8$ for 9 dof for 4U 1608–52 and $\chi^2 = 2.7$ for 4 dof for 4U 1636–53), with the magnitude of the soft time lags increasing marginally faster in 4U 1608–52, $5.0 \pm 0.5 \mu\text{s keV}^{-1}$, than in 4U 1636–53, $3.6 \pm 0.3 \mu\text{s keV}^{-1}$. The lower panel of Fig. 4 shows that in both sources the time lags of the upper kHz QPO are consistent with being zero (the average time lags being $13 \pm 13 \mu\text{s}$ and $4 \pm 3 \mu\text{s}$ for 4U 1608–52 and 4U 1636–53, respectively), independent of energy. (A joint fit to the data of both sources with a linear function, compared to a constant fit, gives an F-test probability of 0.08.)

In the left-hand panel of Fig. 5, we plot separately the intrinsic coherence of the lower and the upper kHz QPO in 4U 1608–52 and 4U 1636–53 using all the frequency intervals combined. For this figure, we calculated the intrinsic coherence of photons with average energy indicated in the x-axis with respect to photons with energies around 3.75 and 4.2 keV for the lower kHz QPO in 4U 1608–52 and 4U 1636–53, respectively, and with respect to photons with energies around ~ 7.5 and 10.2 keV for the upper kHz QPO in 4U 1608–52 and 4U 1636–53, respectively. In both sources, the intrinsic coherence of the lower kHz QPO remains constant at ~ 0.6 (where an intrinsic coherence of 1 indicates perfect correlation between the two signals) up to 14 keV (upper-left panel), and it drops to 0 (indicating that the signals are completely uncorrelated) above that energy. For both sources, the intrinsic coherence of the upper kHz QPO is consistent with being zero across the full energy range (lower-left panel).

Finally, in the right-hand panel of Fig. 5, we plot separately the intrinsic coherence of the lower and the upper kHz QPO in 4U 1608–52 and 4U 1636–53 as a function of the frequency of either

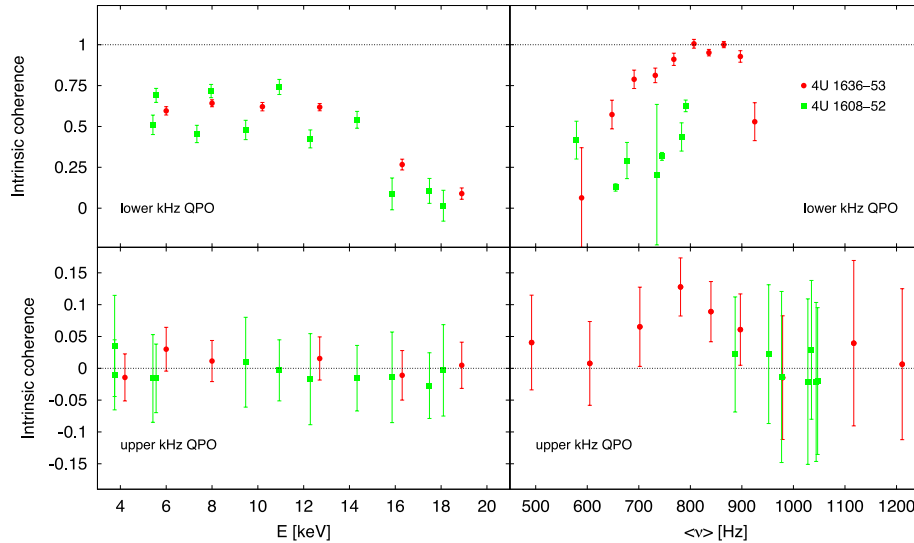


Figure 5. Intrinsic Fourier coherence versus energy (left-hand panels) and versus frequency (right-hand panels) of the lower and the upper kHz QPOs (upper and lower panels, respectively) in 4U 1608–52 (green squares) and 4U 1636–53 (red circles). For the plot in the left-hand panels, we calculated the intrinsic coherence of the QPOs combining all frequency intervals, for photons with energies given by the value on the x -axis relative to photons with energies around 3.75 and ~ 7.5 keV for the lower and the upper kHz QPO in 4U 1608–52, respectively, and 4.2 and 10.2 keV for the lower and the upper kHz QPO in 4U 1636–53, respectively. For the plot in the right-hand panels, we calculated the intrinsic coherence of photons in the 4–12 keV band relative to photons in the 12–20 keV band. The horizontal lines indicate the intrinsic coherence values of 1 and 0 in the upper and the lower panels, respectively.

of the kHz QPO. For this figure, we calculated the intrinsic coherence of photons in the 4–12 keV band relative to photons in the 12–20 keV band. For both sources, the intrinsic coherence of the lower kHz QPO (upper panel) increases from ~ 0 –0.4 at ~ 600 Hz to ~ 1 and ~ 0.6 at ~ 800 Hz in 4U 1636–53 and 4U 1608–52, respectively. In 4U 1636–53, the intrinsic coherence drops significantly above ~ 900 Hz, going down to ~ 0.5 at ~ 920 Hz. In 4U 1608–52, we do not have data above ~ 800 Hz to assess whether, as in the case of 4U 1636–53, the coherence also drops above that frequency. In both sources, the intrinsic coherence of the upper kHz QPO (lower panel) is consistent with being zero over the full frequency range of the upper kHz QPO, except for 4U 1636–53 over a narrow frequency range between ~ 700 and ~ 900 Hz where the intrinsic coherence marginally increases reaching a value of ~ 0.13 at ~ 780 Hz and then it drops again to 0 at ~ 950 Hz. In the case of 4U 1608–52, we do not have data below ~ 880 Hz, and hence we cannot exclude a similar behaviour to that seen in 4U 1636–53.³

4 DISCUSSION

We studied the energy and frequency dependence of the time lags and intrinsic Fourier coherence of the kHz QPOs in the NS LMXBs 4U 1608–52 and 4U 1636–53, using data from the *RXTE*. For the first time we analysed the dependence of the time lags and intrinsic coherence of the kHz QPOs upon the QPO frequency. For both sources, the time lags of the lower kHz QPO are soft (the 4–12 keV photons lag the 12–20 keV ones), and of the order of a few tens of microseconds. In 4U 1636–53, the soft time lags of the lower kHz QPO remain more or less constant at $25 \mu\text{s}$ as the QPO frequency changes from 500 to 850 Hz whereas, when the QPO frequency increases further, the soft lags decrease to $10 \mu\text{s}$. In 4U 1608–52, the soft QPO photons lag the hard ones by $40 \mu\text{s}$, independent

of QPO frequency from 580 up to 780 Hz, the highest frequency reached by this QPO in our data of this source. In both sources, the time lags of the upper kHz QPO are inconsistent with the time lags of the lower kHz QPO. In 4U 1636–53, the hard photons lag the soft ones, so-called hard lags, by $11 \pm 3 \mu\text{s}$, independent of QPO frequency from 500 to 1200 Hz. We found consistent results for the upper kHz QPO in 4U 1608–52, albeit with larger error bars than in 4U 1636–53 over the narrower frequency interval covered by the upper kHz QPO in our data, from 880 to 1040 Hz. The intrinsic coherence (4–12 keV versus 12–20 keV photons) of the lower kHz QPO increases from ~ 0 –0.4 at ~ 600 Hz to 1 and ~ 0.6 at ~ 800 Hz in 4U 1636–53 and 4U 1608–52, respectively. In 4U 1636–53, the intrinsic coherence decreases significantly to 0.5 at 920 Hz, while in 4U 1608–52 we do not have data above 800 Hz. On the contrary, in both sources the intrinsic coherence of the upper kHz QPO is consistent with zero over the full frequency range of the QPO, except in 4U 1636–53 between ~ 700 and ~ 900 Hz where the intrinsic coherence marginally increases to 0.13 at 780 Hz and then drops again to 0 at ~ 950 Hz. Since we do not have data below 880 Hz, we cannot exclude a similar behaviour in the case of 4U 1608–52. We further studied the dependence of the time lags and intrinsic Fourier coherence of the kHz QPOs upon energy. This is the first such study in the case of the time lags of the upper kHz QPO and the intrinsic coherence of both the lower and the upper kHz QPO. We found that in both sources the intrinsic coherence of the lower kHz QPO remains constant at ~ 0.6 between 5 and ~ 12 keV, and drops to zero above that energy. The intrinsic coherence of the upper kHz QPO in both sources is consistent with being zero across the full energy range. We confirmed, with finer energy resolution than in previous analysis (Vaughan et al. 1997, 1998; Kaaret et al. 1999) that in both sources the soft time lags of the lower kHz QPO range from ~ 10 to $\sim 100 \mu\text{s}$, with the soft lags increasing with energy. In both sources, the time lags of the upper kHz QPO are consistent with zero and independent of energy.

Regardless of whether the time lags of the kHz QPOs are soft or hard (negative or positive, respectively, according to the convention

³ We note that the precise value of the intrinsic coherence depends on the value of the noise correction in each band (Vaughan & Nowak 1997).

we used in this paper), light travel time arguments indicate that the size of the region over which these time lags are produced must be smaller than $\sim 3\text{--}30$ km, depending on the energies used to measure the lags (see e.g. Figs 3 and 4). The small change of the time lags with QPO frequency imply that the geometry of the region that produces the lags does not change much either. For instance, the fact that the time lags of the upper kHz QPO in 4U 1636–53 are consistent with being constant at $11 \pm 3 \mu\text{s}$ while the frequency of the QPO changes from ~ 490 to ~ 1210 Hz implies that the size of the region, or the path length, over which the QPO photons are delayed can at most change by ~ 1 km. Note that if the frequency of the upper kHz QPO is the Keplerian frequency at the inner edge of the accretion disc (Miller et al. 1998; Stella & Vietri 1999), for an NS mass of $\sim 1.4\text{--}2.2 M_{\odot}$ the inferred inner-disc radius would change by 12–14 km.

Zoghbi et al. (2010) proposed that the ~ 30 -s soft lags seen in the active galactic nucleus 1H 0707–495 at Fourier frequencies higher than $\sim 5 \times 10^{-4}$ Hz are due to reflection of photons from the corona around the central black hole off the accretion disc. If this scenario is correct and holds also for accreting NS systems, besides photons from the corona one would also have to consider relatively hard photons coming from the NS surface or boundary layer (Sunyaev & Shakura 1986; Kluzniak 1988) reflecting off the accretion disc (Cackett et al. 2010; D’Aì et al. 2010; Sanna et al. 2013). On the other hand, all models of the kHz QPOs propose that the frequency of either the lower or the upper kHz QPO is equal to the Keplerian frequency at the inner edge of the accretion disc (e.g. Miller et al. 1998; Osherovich & Titarchuk 1999; Stella & Vietri 1999). In all these models, changes of the QPO frequency reflect changes of the inner radius of the accretion disc, set by the interplay between mass accretion rate, radiation pressure and radiation drag (Miller et al. 1998; Sanna et al. 2012). A combination of these two ideas (soft lags due to reflection off the accretion disc and one of the kHz QPOs being Keplerian) appears to be consistent (at least qualitatively) with our measurements of the magnitude of the time lags of the lower kHz QPO and their dependence on QPO frequency: In 4U 1636–53, the magnitude of the soft time lags of the lower kHz QPO decreases by $\sim 15 \mu\text{s}$ as the frequency of the QPO increases from ~ 600 to ~ 900 Hz (see Fig. 3); the relative distance between the region in the corona or the NS surface or boundary layer where the hard QPO photons are produced and the region in the disc where the soft QPO photons are reflected would then change by ~ 5 km or less. If we identify the lower kHz QPO as the Keplerian frequency at the inner edge of the disc (Osherovich & Titarchuk 1999), the frequency range spanned by the QPO implies a change of the radius of the inner edge of the disc of $\sim 5\text{--}6$ km for NS masses of $\sim 1.4\text{--}2.2 M_{\odot}$. If instead we identify the upper kHz QPO as the Keplerian one (Miller et al. 1998; Stella & Vietri 1999), using the typical frequency separation of ~ 300 Hz between the two simultaneous kHz QPOs in this source (Wijnands et al. 1997; Méndez et al. 1998c), the frequency range spanned by the lower kHz QPO implies a change of the inner disc radius of ~ 4 km for the same range of NS masses. Not only the magnitude of the soft time lags of the lower kHz QPO and their variation match the typical size and range of inferred inner-disc radii (assuming a localized source of hard photons, e.g. the NS surface), but also the magnitude of the soft time lags decreases as the QPO frequency increases (and the inferred inner-disc radius decreases), which would be the expected behaviour if the bulk of the reflection takes place close to the inner edge of the disc. It is also interesting to note that in the case of 1H 0707–495 the lags become negative at $\sim 5 \times 10^{-4}$ Hz, which scales up to ~ 1200 Hz if the mass ratio of the black hole

in 1H 0707–495 (Zoghbi et al. 2010) and the NS in 4U 1636–53 (Barret et al. 2006) is $\sim 5 \times 10^6 M_{\odot}/2 M_{\odot}$.

The time lags of the upper kHz QPO in 4U 1636–53 are hard, almost 4σ different from zero, and significantly different from the time lags of the lower kHz QPO in this source. The measurements in 4U 1608–52 have larger errors, but they are consistent with those in 4U 1636–53. It appears difficult to explain simultaneously the hard time lags of the upper kHz QPO and the soft time lags of the lower kHz QPO as reflection off the accretion disc around the NS. On one hand, the photons reflected off the disc would have to be harder than the incident photons from the corona or the NS surface or boundary layer in the case of the upper kHz QPO, whereas, in the scenario sketched in the previous paragraph for the lower kHz QPO, the reflected photons were assumed to be softer than the incident ones. If both the soft and hard time lags of, respectively, the lower and the upper kHz QPO are due to reflection off the accretion disc, this would require, for instance, that for the lower kHz QPO the incident photons came from the corona or relatively hot areas of the NS surface or boundary layer and reflected off relatively cold areas in the disc, whereas for the upper kHz QPO the incident photons came from relatively cold areas in the NS surface or boundary layer and reflected off relatively hot areas in the disc. This scenario seems unlikely, given that the disc temperature decreases with distance from the NS and hence from travel time arguments, as the inner disc radius changes, the magnitude of the time lags of the upper kHz QPO would have to change more than that of the lower kHz QPO, contrary to what we observed.

The oscillation mechanism that produces the quasi-periodic variability likely takes place in the disc; it is relatively straightforward to find characteristic (dynamical) frequencies in the accretion disc that match the observed frequencies of the kHz QPOs (e.g. Miller et al. 1998; Osherovich & Titarchuk 1999; Stella & Vietri 1999; Abramowicz et al. 2003), whereas it is more difficult to find a ‘clock’ somewhere else in the accretion flow. However, the emission from the disc alone cannot explain the rms amplitudes of the kHz QPOs, since in some cases the modulated luminosity in the QPOs is ~ 15 per cent, whereas at the same time the emission of the disc is $\lesssim 10$ per cent of the total luminosity of the source. The steep increase of the rms amplitude with energy implies a large modulation of the emitted flux, up to ~ 20 per cent at $\sim 25\text{--}30$ keV (see e.g. Berger et al. 1996), at energies where the contribution of the disc is negligible. Hence, while the oscillation mechanism probably takes place in the disc, the signal is likely modulated and amplified by one of the high-energy spectral components in these sources (Méndez 2006; Sanna et al. 2010). The interpretation that the time lags of the kHz QPO are due to reflection off the accretion disc needs to address the fact that the rms spectrum of the kHz QPOs is hard (Berger et al. 1996; Méndez et al. 2001), and that the frequency-resolved energy spectrum of the kHz QPOs matches the spectrum of the boundary layer in these systems (Gilfanov, Revnivtsev & Molkov 2003). We note that in 4U 1608–52 the rms spectrum of the lower kHz QPO is steeper than that of the upper kHz QPO (Méndez et al. 2001); it is unclear whether this is a general feature of the kHz QPOs, and how the rms spectrum of the kHz QPO depends, for instance, upon the frequency of the kHz QPO itself. This may be a relevant piece of information given the opposite sign of the time lags of the lower and the upper kHz QPO in 4U 1636–53. At any rate, quantitative calculations would be needed to assess whether specific geometries, either of the emitter or the reflector, could explain both the time lags and the rms spectra of both kHz QPOs under the assumption that the time lags of both kHz QPOs are due to reflection off the accretion disc in accreting NS systems.

Other proposed scenarios for producing time lags, be them soft or hard, involve Compton up- or down-scattering (Lee & Miller 1998; Lee et al. 2001; Falanga & Titarchuk 2007). For instance, Falanga & Titarchuk (2007) proposed a model for the time lags of the pulsations of accreting millisecond pulsars in which hard photons emitted from the accreting column or from a hot corona around the NS are Compton down-scattered by cold plasma in the accretion disc or the NS surface. In this model, time lags of the order of a few to a few tens of microseconds imply that the region where the time lags are produced must be small (see below). In principle, the same model would apply to the case of the kHz QPOs, as long as the emission and scattering regions are the same as in the case of millisecond pulsars. This would be the case if the QPO signal comes from the corona or the NS surface or boundary layer, and the down-scattering medium is the atmosphere of the disc. If we fit the energy-dependent time lags of the lower kHz QPO (Fig. 4) with this model, we find number density of the down-scattering medium of $\sim 1\text{--}2 \times 10^{20} \text{ cm}^{-3}$ in both sources, which is about four orders of magnitude smaller than the expected mid-plane density of a standard α -disc extending close to the surface of a $1.7\text{-}M_{\odot}$ and 10-km NS with a spin period of ~ 600 Hz and a magnetic field of $1\text{--}2 \times 10^8$ G (see, for instance, Frank, King & Raine 2002). For an optical depth $\tau = 1\text{--}5$, the size of the down-scattering region would be $\sim 100\text{--}700$ m. The fits with this model are, however, not good, with χ^2 values of 42 for 9 dof and 68 for 4 dof for 4U 1608–52 and 4U 1636–53, respectively. If we apply the same model to the time lags of the upper kHz QPO we find that the number density of the up-scattering medium is $\sim 6 \times 10^{18} \text{ cm}^{-3}$ in both sources, about one order of magnitude lower than in the case of the lower kHz QPO; correspondingly, the size of the up-scattering region in this case would be $\sim 3\text{--}16$ km for $\tau = 1\text{--}5$. While the fits are better than in the case of the lower kHz QPO, they are still formally unacceptable, with χ^2 values of 12 for 9 dof and 6 for 4 dof for 4U 1608–52 and 4U 1636–53, respectively. One way to justify the low densities and, at least in the case of the lower kHz QPOs, the small size of this region would be if the disc is truncated, for instance at the sonic radius or the radius of the ISCO. In that case the radial velocity of the infalling material would increase rapidly there, and the density of the material would decrease sharply.

For the broad-band noise component, between 0.01 and 100 Hz, in the power density spectra of NS and black hole X-ray binaries hard X-rays lag soft X-rays (e.g. Miyamoto et al. 1988; Ford et al. 1999; Nowak et al. 1999). In those cases, the time lags decrease more or less linearly with frequency (the phase lags, equal to the time lags multiplied by the Fourier frequency, are more or less constant as a function of Fourier frequency). Similarly, hard time lags that increase with energy have been observed for the high-frequency [67 Hz (Cui 1999) and 0.6–8 Hz (Reig et al. 2000)] QPOs in the black hole X-ray binary GRS 1915+105. In the case of the low-frequency QPOs, the time lags become soft as the frequency of the QPO increases above ~ 2 Hz.

These broad-band frequency-independent hard phase (time) lags and the time lags of the low- and high-frequency QPOs in black hole systems have been interpreted in terms of a Compton up-scattering medium with a radial density gradient (Hua, Kazanas & Titarchuk 1997; Kazanas, Hua & Titarchuk 1997). The idea behind the Compton up-scattering model is that low-energy photons gain energy in each scattering event. More scatterings imply a higher energy of the emerging photons, but also a longer path length, therefore the higher energy photons emerge later than the lower energy photons produced simultaneously. It is intriguing that the time lags of the broad-band noise component and of the upper kHz

QPO (at least in 4U 1636–53) are hard, whereas the time lags of the lower kHz QPO are soft. As explained by Lee et al. (2001), a Compton down-scattering mechanism for the time lags of the lower kHz QPO is also difficult to reconcile with the observed increase in the rms amplitude of the QPO with energy (Berger et al. 1996), and with the fact that the hard component in the time-averaged spectrum of these systems is interpreted as Compton up-scattering (Payne 1980; Sunyaev & Titarchuk 1980). Lee et al. (2001) proposed a Compton up-scattering model to explain the soft time lags of the lower kHz QPO in which the temperatures of the Comptonizing corona and the source of soft photons oscillate coherently at the QPO frequency (see also Lee & Miller 1998). This could happen, for instance, if the oscillating signal in the soft and hard components are coupled and energy is exchanged between the two. The model of Lee et al. (2001) is able to reproduce both the energy dependence of the time lags and the rms spectrum of the lower kHz QPO in 4U 1608–52, yielding sizes of the corona of $\sim 4\text{--}25$ km. Interestingly, since in this model the soft lags are due to the temperature of the corona and that of the source of soft photons oscillating coherently at the same frequency, one would expect that if the time lags of one of the kHz QPOs are soft, the time lags of the other kHz QPO would not (because the signal of the other kHz QPO has a different frequency, and the temperature of the hard and the soft components cannot oscillate coherently at two different frequencies), similar to what we observed. Under this scenario, the signal of the upper kHz QPO would be up-scattered in the hard component, producing the usual hard lags, as we observed. Also, the fact that time lags of the lower kHz QPO are the ones that are soft implies that the exchange of energy between the corona and the source of soft photons is more strongly linked to the oscillations of the lower than those of the upper kHz QPO. We note that in the case of 1H 0707–495 (Zoghbi et al. 2010), while the high-frequency (above $\sim 5 \times 10^{-4}$ Hz) time lags are soft, the low-frequency broad-band time lags are hard, similar to those in Galactic black hole and NS binaries, and hence the same mechanism may be operating in this case.

The intrinsic coherence of the lower kHz QPO in 4U 1636–53 is consistent with being 0 when the QPO frequency is at ~ 600 Hz and increases more or less linearly to 1 as the QPO frequency increases up to 800 Hz (see upper-right panel of Fig. 5). The intrinsic coherence then remains more or less constant at 1 as the QPO frequency increases further, up to 900 Hz, and then drops abruptly to 0.5 when the QPO frequency is at ~ 920 Hz. The intrinsic coherence of the lower kHz QPO in 4U 1608–52 follows a similar trend, albeit being slightly lower than the one in 4U 1636–53, up to ~ 790 Hz, where it reaches a value of ~ 0.6 (see the same panel of Fig. 5). We do not have data at frequencies higher than this for the lower kHz QPO in 4U 1608–52, and therefore we cannot test whether the intrinsic coherence also drops in this source at high QPO frequencies. The behaviour of the intrinsic coherence with QPO frequency in these two sources is remarkably similar to the one of the quality factor, Q , of the lower kHz QPO with QPO frequency (Barret et al. 2006). The quality factor is defined as the ratio of the centroid frequency to the FWHM of the QPO; this quantity is sometimes also called coherence of the QPO, although it is different from the intrinsic Fourier coherence discussed up to here. To avoid confusion, we will call it here either the quality factor or Q , and we reserve the term coherence to refer to the intrinsic Fourier coherence of the signal. Barret et al. (2006, see also Barret et al. 2005b) proposed that the drop of the quality factor at high QPO frequencies in the lower kHz QPO of 4U 1636–53, 4U 1608–52 and four other sources reflects the fact that the inner edge of the accretion disc reaches the radius

of the ISCO. Méndez (2006, see also Sanna et al. 2010) argued that the drop of Q at high QPO frequencies could be due (at least in part) to changes in the accretion flow around the NS that affects the modulation mechanism that amplifies the signal of the QPO at different energies. The striking similarity between the two curves, on one side the intrinsic coherence of the lower kHz QPO versus QPO frequency (upper right-hand panel of Fig. 5) and on the other side the quality factor of the lower kHz QPO versus QPO frequency (see e.g. fig. 4 of Barret, Olive & Miller 2007), provides a strong indication that the two phenomena are related. We reiterate that, in principle, the two quantities are independent; the quality factor measures the lifetime of the phenomenon that produces the QPO, whereas the intrinsic Fourier coherence measures the degree of correlation between the QPO signal in two different energy bands. While these two quantities can be correlated (as it appears to be the case for these two sources), there is no guarantee that this should always be the case.

An intrinsic coherence of one is equivalent to having two light curves that are related by a linear transform. On the contrary, if the two light curves are related by a non-linear transfer function, power from a given frequency in one light curve is distributed among harmonics of that frequency, leading to a loss of coherence (Bendat & Piersol 1986; Vaughan et al. 1997). The loss of intrinsic coherence both at low and high QPO frequencies in 4U 1636–53 (and at low frequencies in 4U 1608–52) entails that for the lower kHz QPO, the 4–12 keV and 12–20 keV light curves are related by a non-linear transfer function when the QPO frequency is either low or high. An example of such a signal would be that of two light curves produced by local temperature fluctuations such that the average temperature of the emitting material is much smaller than the energy of the emitted photons. In that case, the two light curves are in the tail of the Wien’s spectrum, and they are related to each other via a non-linear transfer function (see Vaughan et al. 1997 for details). Another example of two signals yielding low intrinsic coherence would be that of multiple flaring regions when more than one of those regions contribute to both light curves, even if the individual regions produce perfectly coherent variability (Vaughan et al. 1997).

Zoghbi et al. (2010) found that the intrinsic coherence between the soft and hard light curves, 0.3–1.0 keV and 1.0–4.0 keV, respectively, in 1H 0707–495 are high, $\gtrsim 0.75$, up to a frequency of $\sim 5 \times 10^{-3}$ Hz. Above that frequency the coherence drops because, according to their simulations, the Poisson noise starts to dominate the variability. It is unclear whether reflection off the accretion disc will always produce a high intrinsic coherence, or how the value of the intrinsic coherence depends on properties of the accretion disc, e.g. the size of the inner radius of the disc. On the other hand, in the model of Lee et al. (2001), the temperatures of the corona or boundary layer and the accretion disc oscillate coherently; therefore, the soft and hard light curves are related by a linear transfer function and the intrinsic coherence is high. The size of the Comptonizing medium implied by the fits by Lee et al. (2001) to the data of 4U 1608–52 suggests that a large region of that medium is involved in the nearly coherent variability, hinting at a global mode in the corona or the boundary layer. If this is correct, the drop of the intrinsic coherence at low and high QPO frequencies would be due to the interplay between the excitation and damping mechanisms that produce this mode.

On the other hand, the intrinsic coherence of the upper kHz QPO, both in 4U 1608–52 and 4U 1636–53, are consistent with 0 between 500 and 1200 Hz, except for an ~ 200 Hz interval around 800 Hz in 4U 1636–53 where the intrinsic coherence increases

slightly to 0.13. (We do not have data for the upper kHz QPO of 4U 1608–52 in that frequency range.) In the model of Lee et al. (2001), this could be interpreted as the source of soft photons and the Comptonizing medium not oscillating coherently, which could mean that at those frequencies the source of soft photons and the corona are not (or only weakly) coupled. This could be either because the efficiency of the energy exchange between the source of soft photon on one side and the corona on the other is much higher for producing the lower than the upper kHz QPO, or because the global modes in the system would be more easily excited within a limited frequency range, around 800 Hz. It is interesting that the intrinsic coherence of the upper kHz QPO in 4U 1636–53 appears to increase when the frequency of upper kHz QPO is between ~ 700 and ~ 900 Hz, peaking at a QPO frequency of ~ 800 Hz, the same frequency range over which the intrinsic coherence of the lower kHz QPO in this source increases and peaks. The factor of ~ 10 difference between the intrinsic coherence of the lower and the upper kHz QPO at ~ 800 Hz suggests that the global modes are less efficiently excited, or more efficiently damped, in one case than in the other.

In the case of the lower kHz QPO, the intrinsic coherence is relatively high at energies below ~ 12 keV where it begins to decline, and then it finally reaches 0 at about 18 keV. Sanna et al. (2013) analysed six *XMM-Newton* plus *RXTE* spectra of 4U 1636–53 across different spectral states of the source. They fitted the spectra using a model that accounts for the thermal emission from the accretion disc and the NS surface or boundary layer, the Comptonized emission from a corona, plus reflection off the accretion disc. Sanna et al. (2013) found that in all six observations the Comptonized emission from the corona starts to dominate the spectrum above ~ 10 –15 keV. Below this energy, and down to 1–3 keV, both the NS surface or boundary layer and the corona are more or less equally strong. In the context of the model of Lee et al. (2001), our results suggest that the intrinsic coherence is strong over the energy range in which the boundary layer and the Comptonizing corona are more or less equally strong, and drops as soon as one of these two components no longer contributes to the total emission.

Finally, in the light of the above discussion, one can speculate whether the same mechanism that produces the soft time lags and drives the loss of the intrinsic coherence of the lower kHz QPO would also affect the quality factor, Q , of the QPO. As the temperatures of the source of soft photons and the corona start to oscillate less and less coherently, the feedback between these two components becomes less and less efficient, which we propose would lead to a decrease of the quality factor and the strength of the QPO, similar to what we observe (Barret et al. 2006; Méndez 2006; Sanna et al. 2010). The proposal by Barret et al. (2006) for the drop of Q at high frequencies of the lower kHz QPO involves the inner edge of the accretion disc reaching the ISCO. The drop of Q at low frequencies of the lower kHz QPO is not explained in their model. The mechanism that we propose here would work both at low and high kHz QPO frequencies, and would provide a single explanation to the observed drop of the quality factor (and rms amplitude) of the lower kHz QPOs at both ends of the frequency range. This idea would require further modelling of all the effects involved, which is beyond the scope of this paper.

ACKNOWLEDGEMENTS

The author wish to thank Beike Hiemstra, Diego Altamirano and Guobao Zhang for their help and useful discussions. MGBA and

JEH acknowledge the financial support from CAPES. MGBA is grateful to the Kapteyn Astronomical Institute for their hospitality.

NOTE ADDED IN PROOF

A week after we submitted our paper to MNRAS and posted it on arXiv.org, D. Barret posted a similar paper on arXiv.org (Barret 2013) in which he analyses the time lags of the lower kHz QPO in 4U 1608–52 using a smaller data set than ours. Although the focus of Barret’s and this paper are somewhat different, the results that are common in both papers are consistent with each other.

REFERENCES

- Abramowicz M. A., Karas V., Kluzniak W., Lee W. H., Rebusco P., 2003, *PASJ*, 55, 467
- Altamirano D., van der Klis M., Méndez M., Jonker P., Klein-Wolt M., Lewin W. H. G., 2008, *ApJ*, 685, 436
- Barret D., 2013, *ApJ*, 770, 9
- Barret D., Kluzniak W., Olive J. F., Paltani S., Skinner G. K., 2005a, *MNRAS*, 357, 1288
- Barret D., Olive J.-F., Miller M. C., 2005b, *MNRAS*, 361, 855
- Barret D., Olive J.-F., Miller M. C., 2006, *MNRAS*, 370, 1140
- Barret D., Olive J.-F., Miller M. C., 2007, *MNRAS*, 376, 1139
- Bendat J., Piersol A., 1986, *Random Data: Analysis and Measurement Procedures*. Wiley, New York
- Berger M. et al., 1996, *ApJ*, 469, L13
- Bradt H. V., Rothschild R. E., Swank J. H., 1993, *A&AS*, 97, 355
- Cackett E. M. et al., 2010, *ApJ*, 720, 205
- Cui W., 1999, *ApJ*, 524, L59
- D’Ài A. et al., 2010, *A&A*, 516, A36
- Di Salvo T., Méndez M., van der Klis M., 2003, *A&A*, 406, 177
- Falanga M., Titarchuk L., 2007, *ApJ*, 661, 1084
- Ford E. C., van der Klis M., Méndez M., van Paradijs J., Kaaret P., 1999, *ApJ*, 512, L31
- Frank J., King A., Raine D., 2002, *Accretion Power in Astrophysics*, 3rd edn, Cambridge Univ. Press, Cambridge
- Gilfanov M., Revnivtsev M., Molkov S., 2003, *A&A*, 410, 217
- Hua X. M., Kazanas D., Titarchuk L., 1997, *ApJ*, 482, L57
- Jahoda K., Markwardt C. B., Radeva Y., Rots A. H., Stark M. J., Swank J. H., Strohmayers T. E., Zhang W., 2006, *ApJS*, 163, 401
- Kaaret P., Piraino S., Ford E., Santangelo A., 1999, *ApJ*, 514, L31
- Kazanas D., Hua X. M., Titarchuk L., 1997, *ApJ*, 480, 735
- Kluzniak W., 1988, PhD thesis, Stanford University
- Lee H. C., Miller G. S., 1998, *MNRAS*, 299, 479
- Lee H. C., Misra R., Taam R. E., 2001, *ApJ*, 549, L229
- Lochner J., Roussel-Dupré D., 1994, *ApJ*, 435, 840
- Méndez M., 2006, *MNRAS*, 371, 1925
- Méndez M. et al., 1998a, *ApJ*, 494, L65
- Méndez M., van der Klis M., Wijnands R., Ford E. C., van Paradijs J., Vaughan B. A., 1998b, *ApJ*, 505, L23
- Méndez M., van der Klis M., van Paradijs J., 1998c, *ApJ*, 506, L117
- Méndez M., van der Klis M., Ford E. C., 2001, *ApJ*, 561, 1016
- Miller M. C., Lamb F. K., Psaltis D., 1998, *ApJ*, 508, 791
- Miyamoto S., Kitamoto S., Mitsuda K., Dotani T., 1988, *Nat*, 336, 450
- Nowak M. A., Vaughan B. A., Wilms J., Dove J. B., Begelman M. C., 1999, *ApJ*, 510, 874
- Osherovich V., Titarchuk L., 1999, *ApJ*, 522, L113
- Payne D. G., 1980, *ApJ*, 237, 951
- Psaltis D., Belloni T., van der Klis M., 1999, *ApJ*, 520, 262
- Reig P., Belloni T., van der Klis M., Méndez M., Kylafis N. D., Ford E. C., 2000, *ApJ*, 541, 883
- Sanna A. et al., 2010, *MNRAS*, 408, 622
- Sanna A., Méndez M., Belloni T., Altamirano D., 2012, *MNRAS*, 424, 2936
- Sanna A., Hiemstra B., Méndez M., Altamirano D., Belloni T., Linares M., 2013, *MNRAS*, 432, 1144
- Stella L., Vietri M., 1999, *Phys. Rev. Lett.*, 82, 17
- Sunyaev R. A., Shakura N. I., 1986, *SvA*, 12, L117
- Sunyaev R. A., Titarchuk L. G., 1980, *A&A*, 86, 121
- van der Klis M., 2006, in Lewin W., van der Klis M., eds, *Compact Stellar X-ray Sources*. Cambridge Univ. Press, Cambridge, p. 39
- van Straaten S., van der Klis M., di Salvo, Tiziana, Belloni T., 2002, *ApJ*, 568, 912
- van Straaten S., van der Klis M., Méndez M., 2003, *ApJ*, 596, 1155
- Vaughan B. A., Nowak M. A., 1997, *ApJ*, 474, L43
- Vaughan B. A. et al., 1997, *ApJ*, 483, L115
- Vaughan B. A. et al., 1998, *ApJ*, 509, L145
- Wijnands R. A. D., van der Klis M., van Paradijs J., Lewin W. H. G., Lamb F. K., Vaughan B., Kuulkers E., 1997, *ApJ*, 479, L141
- Zoghbi A., Fabian A. C., Uttley P., Miniutti G., Gallo L. C., Reynolds C. S., Miller J. M., Ponti G., 2010, *MNRAS*, 401, 2419
- Zoghbi A., Uttley P., Fabian A. C., 2011, *MNRAS*, 412, 59

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.