

University of Groningen

Emission characteristics of water in the Universe

Poelman, Dieter Roel

IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.

Document Version

Publisher's PDF, also known as Version of record

Publication date:

2007

[Link to publication in University of Groningen/UMCG research database](#)

Citation for published version (APA):

Poelman, D. R. (2007). *Emission characteristics of water in the Universe*. s.n.

Copyright

Other than for strictly personal use, it is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), unless the work is under an open content license (like Creative Commons).

The publication may also be distributed here under the terms of Article 25fa of the Dutch Copyright Act, indicated by the "Taverne" license. More information can be found on the University of Groningen website: <https://www.rug.nl/library/open-access/self-archiving-pure/taverne-amendment>.

Take-down policy

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

Downloaded from the University of Groningen/UMCG research database (Pure): <http://www.rug.nl/research/portal>. For technical reasons the number of authors shown on this cover page is limited to 10 maximum.

Chapter 5

Water line emission from the high-mass star-forming region AFGL 2591

Observations of water lines are a sensitive probe of the geometry, dynamics and chemical structure of dense molecular gas. We investigate the diagnostic value of specific water transitions in high-mass star-forming regions. As a test case, we apply our models to the AFGL 2591 region. Similarities and differences of constant and jump abundance models are displayed, as well as when an outflow is incorporated. In general, for models with a constant water abundance, the ground-state lines, i.e., $1_{10}\text{-}1_{01}$, $1_{11}\text{-}0_{00}$, and $2_{12}\text{-}1_{01}$, are predicted in absorption, all the others in emission. This behaviour changes for models with a water abundance jump profile in that the line profiles for jumps by a factor of $\sim 10\text{--}100$ are similar to the line shapes in the constant abundance models, whereas larger jumps lead to emission profiles. Asymmetric line profiles are found for models with a cavity outflow and depend on the inclination angle. Models with an outflow cavity are favoured to reproduce the SWAS observations of the $1_{10}\text{-}1_{01}$ ground-state transition. PACS spectra will tell us about the geometry of these regions, both through the continuum and through the lines. It is found that the low-lying transitions of water are sensitive to outflow features, and represent the excitation conditions in the outer regions. High-lying transitions are more sensitive to the adopted density and temperature distribution which probe the inner excitation conditions. The Herschel mission will thus be very helpful to constrain the physical and chemical structure of high-mass star-forming regions such as AFGL 2591.

Based on Poelman, D. R., & van der Tak, F. F. S., A&A submitted

5.1 Introduction

High-mass stars, more than their low-mass counterparts, play a critical role in the energetics and dynamical feedback into the interstellar medium (ISM). Despite this, the formation mechanisms of massive stars are still a puzzle, owing to their occurrence in more distant crowded stellar clusters, shorter formation timescales, and formation in regions of high visual extinction. Despite observational obstructions, promising theoretical work has been done on the formation of high-mass stars (e.g., Bonnell et al. 1998; Klessen 2001; McKee & Tan 2002, 2003; Klessen et al. 2007). Whereas low- and intermediate-mass stars ($M \lesssim 10 M_{\odot}$) are formed through gravitational collapse and subsequent disk accretion (Palla & Stahler 1993), radiation pressure on the dust in the infalling gas prohibits high-mass stars to sufficiently accrete gas and thus limits the stellar mass at 20–40 M_{\odot} (Yorke & Krügel 1977; Wolfire & Cassinelli 1987; Beech & Mitalas 1994). However, observations show that considerably more massive stars exist. Two formation scenarios have been proposed to work out this contradiction: non-spherical accretion (Yorke & Sonnhalter 2002) and merging of lower mass stars (Bonnell et al. 1998; Bonnell & Bate 2002). In order to distinguish between these two scenarios detailed observations, sensitive to a range of size scales and physical conditions, are required. For this, interferometers can provide the spatial distribution of dust and gas at the smallest size scales, while single dish telescopes equipped with large format, wide bandwidth, detectors are capable of simultaneously imaging a large number of molecules (albeit at lower spatial resolution).

Up to date, there is growing evidence that disks are present around massive young stars as reports of direct imaging of disks in regions of high-mass star formation are increasing in number (e.g., Beltrán et al. 2004; Zhang et al. 2002; Sandell et al. 2003; van der Tak & Menten 2005; van der Tak et al. 2006). In addition to direct imaging of disk-like structures, outflows provide indirect evidence for disks around massive young stars. In particular, collimated outflows are a sign of accretion (Beuther et al. 2002). Outflows are well known phenomena accompanied with sites of low-mass star formation, transporting angular momentum away from the star. The situation in the formation of high-mass stars is less clear, since spatial resolution has been lacking to resolve the outflows. However, over the last years, outflows from high-mass stars have received increasing attention (see recent reviews by, e.g., Shepherd 2003, 2005; Cesaroni 2005; Arce et al. 2007). It is found that the higher-mass objects appear to continue smoothly the correlation found in T Tauri stars between outflow and accretion signatures. These results suggest that the disk-outflow connection found in low-mass pre-main-sequence stars extends to more massive objects, and favour the accretion scenario. Besides this, Krumholz et al. (2005) show that outflows around massive stars lead to a significant anisotropy in the

stellar radiation field, hence reducing the radiation pressure experienced by the gas in the infalling envelope. As a result, a larger fraction of the material in the envelope can be used in the formation process of a massive star as opposed to models without outflows.

The dynamical and physical evolution during star formation is accompanied by strong chemical evolution. The cold-core phase, dominated by relatively simple molecular species, is transformed into a hotbed of more complex molecules after the protostellar object heats up the surrounding gas. Some of these molecules, e.g., H_2 , CO , and H_2O , play an important role in regulating the temperature, hence pressure, via the process of heating and cooling through line absorption and emission. Doty & Neufeld (1997) constructed models for the thermal balance, chemistry, continuum and line radiative transfer in dense molecular cloud cores with an embedded protostar. They find that due to temperature and density gradients, molecular abundances are not constant within dense molecular cloud cores, i.e., freeze-out of water onto dust grains for temperatures below the sublimation temperature of ~ 100 K and a high abundance of water in regions where the temperature is larger than 100 K. Also, this behaviour has been found by Ceccarelli et al. (2000) who modeled the structure of the collapsing envelope around the low-mass protostar IRAS 16293-2422. Besides this, outflows not only affect the physical state of the interstellar gas, but are also accompanied with chemical changes. Observational estimates of molecular abundances in outflow regions have suggested that particular species, e.g., SiO , CH_3OH , HCN , H_2CO , SO , and SO_2 are significantly enhanced (e.g., Bachiller & Perez Gutierrez 1997). In particular for water, the elevated gas temperatures found in the outflows increase the rate coefficients for the neutral-neutral reactions, leading to an enhanced water production rate. Consequently, the abundance of water shows significant variations up to a few orders of magnitude from one region to another. This large variation in abundance makes water a powerful diagnostic of the physical structure of the region and of the fundamental chemical interactions between the gas and the grains. Observations, together with theoretical modeling of the available rotational and ro-vibrational lines from species and their isotopes, are necessary to shed light on the density, temperature, and chemical structure over a wide range of conditions relevant for regions of star formation. Water in particular, with its many rotational lines, is very well suited to study the physical conditions in regions of low-mass (van Kempen et al. 2007) and high-mass (this chapter) star formation.

In this chapter, we model the excitation of water in the high-mass star-forming region AFGL 2591. The aim of our research is to make predictions for the line strengths and profiles of various ortho- and para- H_2O (o- and p- H_2O) transitions to be observed with the *Heterodyne Instrument for the Far Infrared* (HIFI) and the *Photodetector Array Camera and Spectrometer* (PACS) on board of Herschel, and to study how the results depend on the

adopted geometry and chemical structure. Resulting line strengths for specific transitions are then compared to existing *Infrared Space Observatory* (ISO) and *Submillimeter Wave Astronomy Satellite* (SWAS) observations.

5.2 The high-mass star-forming region AFGL 2591

While most massive stars form in clusters, GL 2591, located in the Cygnus X region, provides one of the rare cases of a massive star forming in relative isolation. This allows to study the temperature, density, and velocity structure of the circumstellar envelope without confusion from nearby objects. Large columns of dust and gas toward GL 2591 block our view of the stellar photosphere, making it invisible at optical wavelengths, but results in bright infrared emission. The infrared source is associated with a weak radio continuum source (Campbell 1984; Trinidad et al. 2003) and with a bipolar molecular outflow larger than $1'$ in extent (Lada et al. 1984; Mitchell et al. 1992; Hasegawa & Mitchell 1995). Observations of the molecular cloud complexes in the Cygnus X region indicate that most objects are located at the same distance of 1.7 kpc (Schneider et al. 2006). However, in the following we adopt a distance to the Cygnus X region of 1 kpc, fixing the luminosity at $\sim 2 \times 10^4 L_{\odot}$, to allow comparison with previous work (van der Tak et al. 1999). The results presented here depend weakly on distance, making the adopted distance inconsequential.

AFGL 2591 has been observed in water lines over a range of excitation conditions. Helmich et al. (1996) report the detection of more than 30 lines within the bending vibration of water at $6 \mu\text{m}$ using the *Short Wavelength Spectrometer* (SWS, de Graauw et al. 1996) on board ISO (Kessler et al. 1996). van Dishoeck & Helmich (1996) and Boonman & van Dishoeck (2003) derive from observed H_2O spectra around $6 \mu\text{m}$ water abundances of $\sim 10^{-5}$ up to 10^{-4} in hot-core regions. Snell et al. (2000) observed with the *Submillimeter Wave Astronomy Satellite* (SWAS) the ground-state transition of o- H_2O , thereby detecting the H_2O emission from the more extended, cold gas. They find that the H_2O abundances estimated for the hot core gas is at least 100 times larger than in the gas probed by SWAS, the latter being 6×10^{-10} to 10^{-8} . Boonman et al. (2003) observed with the Short and Long Wavelength Spectrometer (SWS and LWS) on board of ISO and with SWAS H_2O lines toward deeply embedded massive protostars, and find that ice evaporation in the warmer envelope and freeze-out in the cold outer parts together with pure gas-phase chemistry reproduces the H_2O observations. This conclusion is strengthened by ground-based observations of the H_2^{18}O isotope (van der Tak et al. 2006).

Despite previous observations of the region, the future launch of ESA's Herschel Space Observatory (Pilbratt 2005), with on board HIFI will provide key information on the physical and chemical conditions in molecu-

lar clouds by means of observing many water lines with higher sensitivity. Moreover, its high angular resolution allows to separate regions. HIFI, a heterodyne receiver, is designed to provide continuous frequency coverage from 480 to 1250 GHz (Band 1–5), while Band 6 will cover the 1410–1910 GHz frequency interval. Going up in frequency, the beam size decreases from 39" to 13", and together with a spectral resolution of 0.2–0.4 km s⁻¹ this instrument will be a powerful tool to probe the kinematics, i.e., infall and/or outflow, and chemical complexity of the AFGL 2591 region. In addition, PACS will take spectral images over the spectral band from 57 to 210 μm. The main advantages of this instrument are the 9" resolution together with a 5 × 5 pixel imaging capability. Although the PACS spectral resolution of $\lambda/\Delta\lambda \sim 1500$ is insufficient to probe the kinematics in regions such as AFGL 2591, mapping of the H₂O and continuum emission will provide supplementary information over a wavelength range inaccessible to HIFI.

5.3 Model set-up

The temperature and density structure of the gas and dust in the envelope surrounding the young stellar object (YSO) AFGL 2591 has been determined by van der Tak et al. (1999, 2000). In short, the resulting profiles are based on a study of single-dish submillimeter spectroscopy, combined with molecular line emission at (sub)-millimeter wavelengths. The temperature structure of the dust was modeled with the dust radiative transfer program of Egan et al. (1988), solving for the thermal balance of the grains as a function of distance to the star. The resulting temperatures follow a $r^{-0.4}$ profile in the outer parts of the envelope, i.e., distances greater than $2-3 \times 10^3$ AU from the star, whereas in the inner envelope the temperature gradient is steeper than $r^{-0.4}$. The density structure is then obtained with the use of the RATRAN code (Hogerheijde & van der Tak 2000) to model the observed molecular lines, i.e., CS, HCO⁺, HCN, and H₂CO, using the temperature distribution calculated by the dust code. It is found that the density follows a power law of the form $n = n_0(r/r_0)^{-\alpha}$, with $\alpha \sim 1-2$. Dust emission profiles are considered as well. In the following, results are presented for $\alpha = 1$ and implications of an adopted $\alpha = 2$ profile are elaborated on in Section 5.5.

Throughout this chapter, a two-dimensional set-up is chosen in which the radiative transfer models consist of 78 logarithmically spaced shells with an outer radius of $\sim 50\,000$ AU. A constant Doppler parameter (1/e width) of 2 km s⁻¹ is assumed for all the models. Fig. 5.1 plots the adopted temperature and density structure as a function of radius for the models with $\alpha = 1$ and 2. It is seen that the $\alpha = 1$ models tend to have a lower density and temperature towards the center of the cloud. The gas and dust tempe-

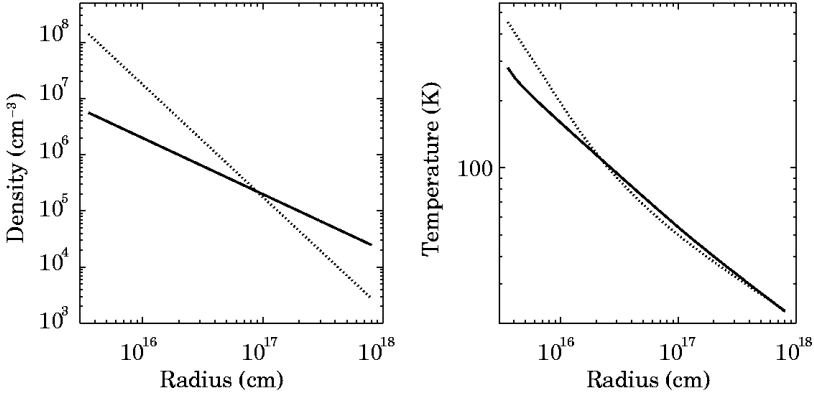


Figure 5.1: The distribution of density (*left*) and temperature of gas and dust (*right*) as a function of radius in the cloud. Solid lines mean $\alpha = 1$, dotted $\alpha = 2$.

temperature is assumed to be coupled, due to the high densities reached throughout most of the cloud. The calculated level populations at each position in the cloud are then used to compute velocity profiles of various transitions of o- and p-H₂O using the program SKY that calculates the sky brightness distribution, part of the RATRAN code. SKY has been tested against other codes. For transitions with a rest frequency within the frequency coverage of HIFI, the resulting sky brightness distribution is convolved with the appropriate beam (depending on the frequency of the transition). In addition, predictions are made for the spectrum to be observed with PACS. Line strengths and profiles are compared with existing ISO observations (Boonman et al. 2003) and therefore convolved with the appropriate ISO-LWS beam, between 66'' and 78'' in diameter, depending on the frequency.

Due to the complex structure and spectroscopic properties of the water molecule, inverted transitions (*masers*) are likely to occur. Maser action, a process of intense radiation confined within small angular sizes with beaming angles in the range of 10^{-1} – 10^{-2} rad (Elitzur 1992), necessitates higher angular resolution than achieved with our adopted multi-zone escape probability method (see Chapter 2, Section 2.3). Moreover, a large negative optical depth, τ , magnifies the $\exp(-\tau)$ term in the escape probability which can hamper convergence. Therefore, we have limited maser growth to a negative optical depth of -0.1 . This has no consequences on the calculated line intensities since non-saturated masers have negligible effect on the remaining level populations. Additional tests have been performed with an optical depth lower limit of -0.08 and -0.12 in order to investigate the effects of the choice of this lower limit. We find that these lower limits

give the same level populations as calculated with a lower limit of -0.1 . Nevertheless, in this chapter, maser action is suppressed, and therefore no predictions are made for maser transitions, e.g., at 22 GHz and 183 GHz.

5.4 Results

5.4.1 Constant water abundance

As a first step, we assume the H_2O abundance to be uniformly distributed. Snell et al. (2000) find, using SWAS observations of the $1_{10}\text{-}1_{01}$ ground-state transition at 557 GHz, that the abundance of o- H_2O relative to H_2 in giant molecular cloud cores varies between $\sim 10^{-10}$ and 10^{-8} . In particular for AFGL 2591, a value of 6×10^{-9} is found. In view of this large observed spread, a constant H_2O abundance with respect to H_2 , ranging from 10^{-10} to 10^{-7} is chosen. Resulting line profiles are plotted in Fig. 5.2, integrated brightness temperatures are listed in Table 5.1. Beside H_2^{16}O lines, multiple H_2^{17}O and H_2^{18}O transitions will be observed with HIFI and PACS for high-mass YSOs. Hence, model results in the case $X(\text{H}_2\text{O})$ is 10^{-10} and 10^{-9} are representative for H_2^{17}O and H_2^{18}O , respectively.

Table 5.1: Integrated intensities, $\int T_{\text{MB}} \Delta\nu$ [K km s^{-1}], for models with a constant H_2O abundance.

Transition	Frequency [GHz]	Water abundance			
		10^{-10}	10^{-9}	10^{-8}	10^{-7}
ortho- H_2O transitions					
$1_{10}\text{-}1_{01}$	557.936	-1.1(-2)	-0.2	-1.2	-1.4
$3_{12}\text{-}3_{03}$	1097.35	3.3(-2)	0.4	4.1	21.8
$3_{12}\text{-}2_{21}$	1153.12	6.0(-3)	7.2(-2)	0.9	12.0
$3_{21}\text{-}3_{12}$	1162.91	4.1(-2)	0.5	5.1	34.6
$2_{21}\text{-}2_{12}$	1661.01	2.0(-3)	0.1	1.1	0.6
$2_{12}\text{-}1_{01}$	1669.90	-0.8	-6.7	-17.7	-24.0
$3_{03}\text{-}2_{12}$	1716.77	6.4(-2)	0.7	4.6	5.0
$3_{30}\text{-}3_{21}$	2196.34	1.1(-4)	1.4(-3)	1.6(-2)	0.2
$4_{14}\text{-}3_{03}$	2640.47	6.6(-4)	7.1(-3)	5.2(-2)	0.1
$2_{21}\text{-}1_{10}$	2773.97	7.9(-4)	9.6(-3)	-2.0(-3)	-0.7
para- H_2O transitions					
$2_{11}\text{-}2_{02}$	752.029	2.1(-2)	0.3	3.0	19.1
$2_{02}\text{-}1_{11}$	987.924	3.2(-2)	0.4	4.3	20.7
$1_{11}\text{-}0_{00}$	1113.34	-0.2	-2.0	-7.1	-10.1
$2_{20}\text{-}2_{11}$	1228.80	2.0(-2)	0.2	2.8	18.7

NOTE: $a(b)$ means $a \times 10^b$

In general, the ground-state lines, i.e., $1_{10}\text{-}1_{01}$, $1_{11}\text{-}0_{00}$ and $2_{12}\text{-}1_{01}$, appear in absorption. All other lines appear in emission. The optical depths in the $1_{10}\text{-}1_{01}$, $1_{11}\text{-}0_{00}$ and $2_{12}\text{-}1_{01}$ transitions are factors of $\gtrsim 10$ higher than the optical depths in the other lines, causing self-absorbed line profiles. When increasing the abundance by factors of 10, the line shapes transform from single-peaked emission into double-peaked self-absorbed profiles for the $2_{02}\text{-}1_{11}$, $2_{21}\text{-}2_{12}$, and $3_{03}\text{-}2_{12}$ transitions. The strongest transitions are found to be those arising from the warm inner region, e.g., $3_{12}\text{-}3_{03}$, $3_{21}\text{-}3_{12}$, and $3_{03}\text{-}2_{12}$.

5.4.2 Water abundance jump profile

In this section we calculate the excitation of water, assuming that the H_2O distribution follows a step function, i.e., a low abundance in the outer regions and a higher abundance in the inner zones. Doty et al. (2006) have constructed models for the chemical evolution in a collapsing massive YSO. These models include a realistic evolution of the central source, as well as ad/desorption of ices from grain mantles as the grains fall in from the cool exterior into the warm interior. They find that this approach reproduces the step-function water distribution inferred observationally (van der Tak et al. 2006). For this reason, models with a water abundance jump profile are computed. The position in the cloud where the jump occurs corresponds to the position where the sublimation temperature of water is reached, assumed to be 100 K in our models. The water abundance outside the sublimation radius is kept constant at 10^{-9} , whereas for regions with temperatures exceeding 100 K, the water abundance is increased from 10^{-8} to 10^{-4} for the different models. Resulting line profiles are plotted in Fig. 5.3, integrated brightness temperatures are listed in Table 5.2.

It is seen that the line profiles of the ground-state lines, i.e., $1_{10}\text{-}1_{01}$, $2_{12}\text{-}1_{01}$, and $1_{11}\text{-}0_{00}$, are self-absorbed, similar to the constant abundance models. Line strengths and profiles of the ground-state transitions of o- and p- H_2O , i.e., $1_{10}\text{-}1_{01}$, $2_{12}\text{-}1_{01}$ and $1_{11}\text{-}0_{00}$, for models with a jump of a factor of 10–100 are equivalent to the line strengths and profiles in case of the model with a constant water abundance of 10^{-9} . Larger jumps, however, do influence the line strengths and profiles for these transitions. Thus, the ground-state lines are partly influenced by the warm inner component, and partly by the cold outer layer. Higher-lying transitions differ from the constant abundance model in that larger line strengths are found even for moderate jumps. This is due to the enhanced inner warm water abundance, the region that favours the excitation of the high-lying transitions.

The peculiar appearance of the line profiles can be explained as a superposition of two constant models, i.e., a small, warm and abundant component together with a cold, large and low abundant component. In other words, a peak or drop in the line profile corresponds to a peak or drop in

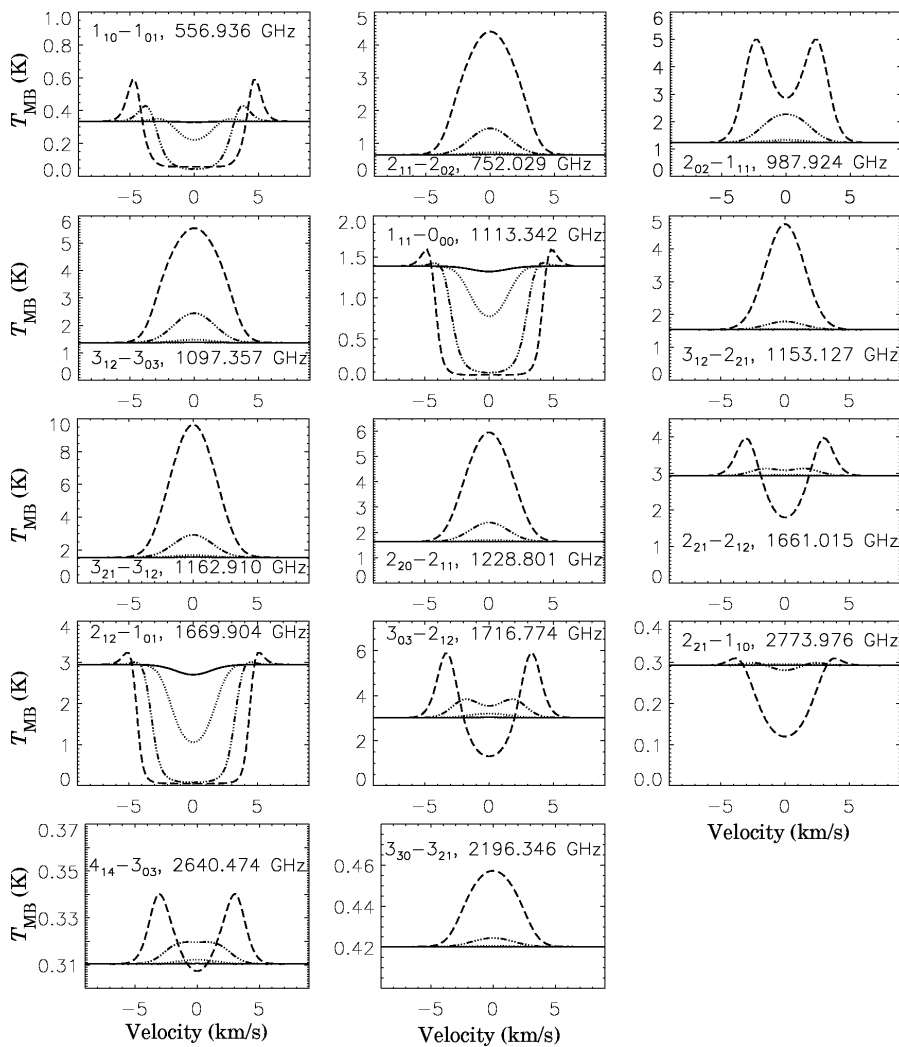


Figure 5.2: Line profiles of ortho- and para-H₂O transitions for models with $\alpha = 1$. The water abundance, i.e., o-H₂O + p-H₂O, is uniformly distributed across the region and ranges from 10^{-10} to 10^{-7} , with 10^{-10} (solid), 10^{-9} (dotted), 10^{-8} (dash-dotted), 10^{-7} (dashed).

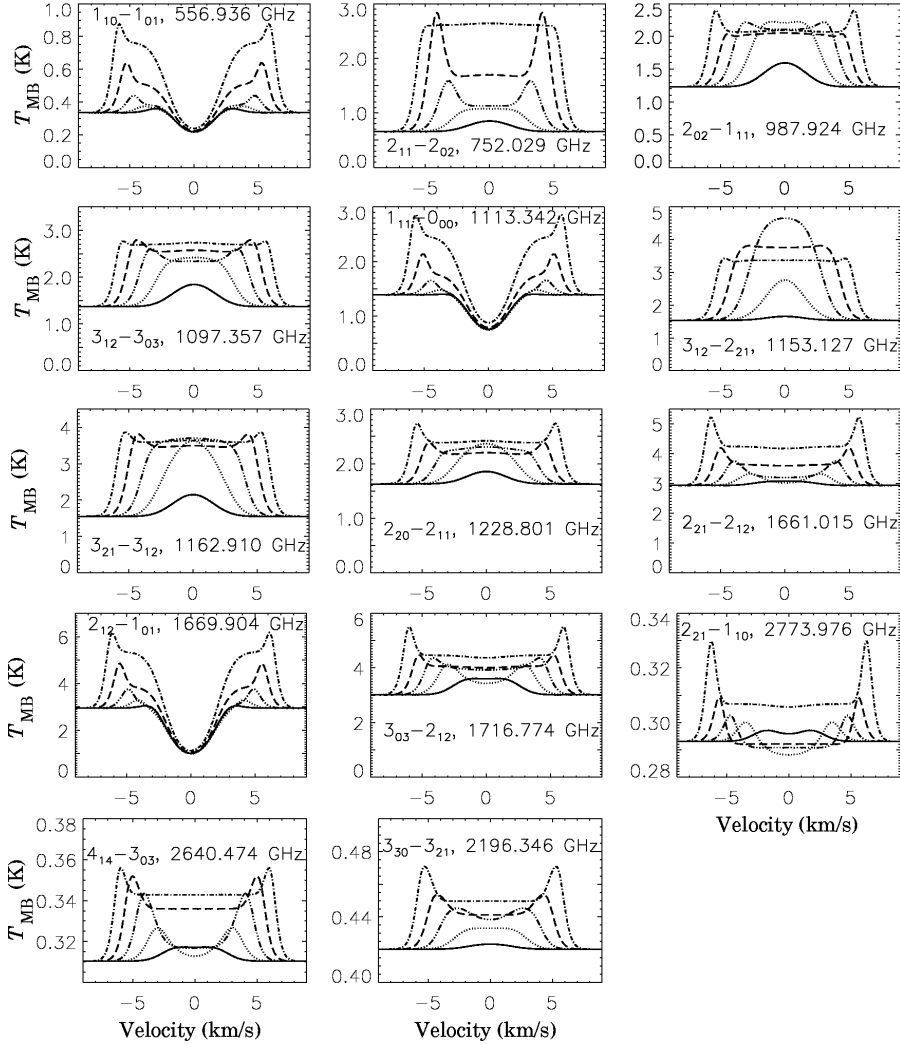


Figure 5.3: Line profiles of H_2O for the jump abundance models with $\alpha = 1$. The water abundances for positions where $T < 100$ K is 10^{-9} , whereas for $T > 100$ K the abundances range from 10^{-8} (solid), 10^{-7} (dotted), 10^{-6} (dash-dotted), 10^{-5} (dashed), 10^{-4} (solid).

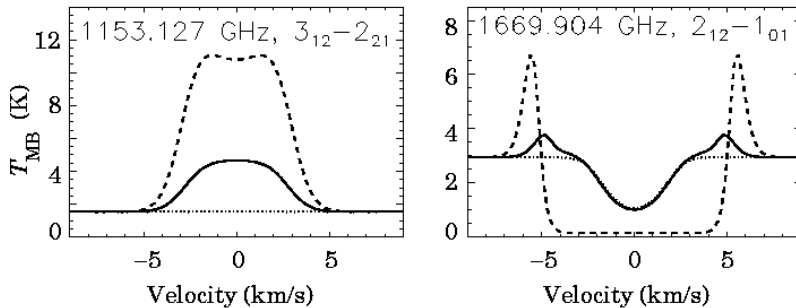


Figure 5.4: Line profiles of two ortho-H₂O transitions. The solid line represents the line profile for a jump with an abundance of 10^{-9} where $T < 100$ K and 10^{-6} where $T > 100$ K. The dotted and dashed line profiles represent the constant water abundance of 10^{-9} and 10^{-6} , respectively.

one of the transitions, when the contributing lines are optically thin. However, the line profile when one or both of the transitions are optically thick is not a simple superposition. In this case, photons coming from line center are lost and only those photons that escape in the line wings will contribute to the final line profile. This can be seen in Fig. 5.4 where we plot, for two different transitions, the resulting jump profiles with an inner water abundance of 10^{-6} , together with the constant water abundance profiles of 10^{-9} and 10^{-6} . It is seen that for the $2_{12}-1_{01}$ transition, the contribution from the optically thick inner 10^{-6} part is limited to those photons escaping in the line wings. However, the resulting jump profile is dominated by the optically thin outer 10^{-9} component. On the contrary, the jump profile for the $3_{12}-2_{21}$ transition is most heavily influenced by the inner component.

One can wonder whether HIFI's future high spectral resolution observations can distinguish between a constant or a jump profile water distribution. Since high-lying transitions, e.g., $3_{12}-3_{03}$, $3_{21}-3_{12}$, $2_{21}-2_{12}$, more than the ground-state lines, are influenced by the warm inner water abundance, these lines are favoured to examine the occurrence of an abundance jump. HIFI's high sensitivity allows to probe features in the line profile of ~ 1 K in Band 6, down to $\sim 0.05-0.1$ K in Band 1 with sufficient integration time. Despite the differences in line strengths, a jump of a factor of $\sim 10-100$ leads to line shapes similar to the profiles in the constant abundance models. Hence, we do not expect HIFI to be able to distinguish between the two scenarios for jumps up to a factor of ~ 100 . However, jumps of a factor of ≥ 1000 lead to line profiles more flattened around line center, with peaks in the line wings. This feature is not seen in the constant models for abun-

Table 5.2: Integrated intensities, $\int T_{\text{MB}} \Delta v$ [K km s^{-1}], for models with a H_2O abundance jump profile. The water abundance outside the sublimation radius is kept constant at 10^{-9} .

Transition	Frequency [GHz]	Inner water abundance				
		10^{-8}	10^{-7}	10^{-6}	10^{-5}	10^{-4}
ortho- H_2O transitions						
1 ₁₀ -1 ₀₁	557.936	-0.2	-0.2	0.1	1.2	3.5
3 ₁₂ -3 ₀₃	1097.35	1.9	6.4	9.6	13.2	17.0
3 ₁₂ -2 ₂₁	1153.12	0.4	4.5	17.3	19.8	20.5
3 ₂₁ -3 ₁₂	1162.91	2.3	9.9	17.0	21.2	25.9
2 ₂₁ -2 ₁₂	1661.01	0.7	2.0	4.6	9.3	18.7
2 ₁₂ -1 ₀₁	1669.90	-6.6	-6.3	-4.5	1.7	15.3
3 ₀₃ -2 ₁₂	1716.77	3.3	6.3	10.8	13.2	21.1
3 ₃₀ -3 ₂₁	2196.34	1.3(-2)	7.7(-2)	0.2	0.3	0.4
4 ₁₄ -3 ₀₃	2640.47	3.9(-2)	8.1(-2)	0.2	0.3	0.5
2 ₂₁ -1 ₁₀	2773.76	2.3(-2)	6.0(-3)	9.0(-3)	2.7(-2)	0.2
para- H_2O transitions						
2 ₁₁ -2 ₀₂	752.029	0.8	2.5	5.5	13.4	21.3
2 ₀₂ -1 ₁₁	987.924	1.4	5.7	7.8	8.6	12.5
1 ₁₁ -0 ₀₀	1113.34	-1.9	-1.9	-1.2	1.5	7.4
2 ₂₀ -2 ₁₁	1228.80	0.8	3.5	5.3	6.5	11.2

NOTE: $a(b)$ means $a \times 10^b$

dances up to 10^{-7} for the high-lying transitions, and thus can be used to determine the water abundance structure.

5.4.3 Cavity

We next consider cylindrically symmetric models in which an outflow cavity is incorporated. Observations confirm a powerful bipolar molecular outflow associated with AFGL 2591 (e.g., Bally & Lada 1983; Lada et al. 1984), ~ 1.5 pc in extent. In addition, van der Tak et al. (2006) constructed a representation of the geometry of AFGL 2591, based on millimeter interferometric observations. They conclude that a young protostar is observed with a disk close to face-on (inclination between 26–38 degrees). Therefore, models are constructed to study the effect on the line strengths and profiles of a low-density outflow with an half-opening angle θ , as function of inclination angle. For this, a model with an outflow cavity with half-opening angle $\theta = 15^\circ$ is computed. Models with larger opening angles are also computed and discussed in Section 5.5. The water abundance outside the cone is set at a constant value of 10^{-9} , and the density and temperature structure follow Fig. 5.1. The outflow is characterised as a low density region with

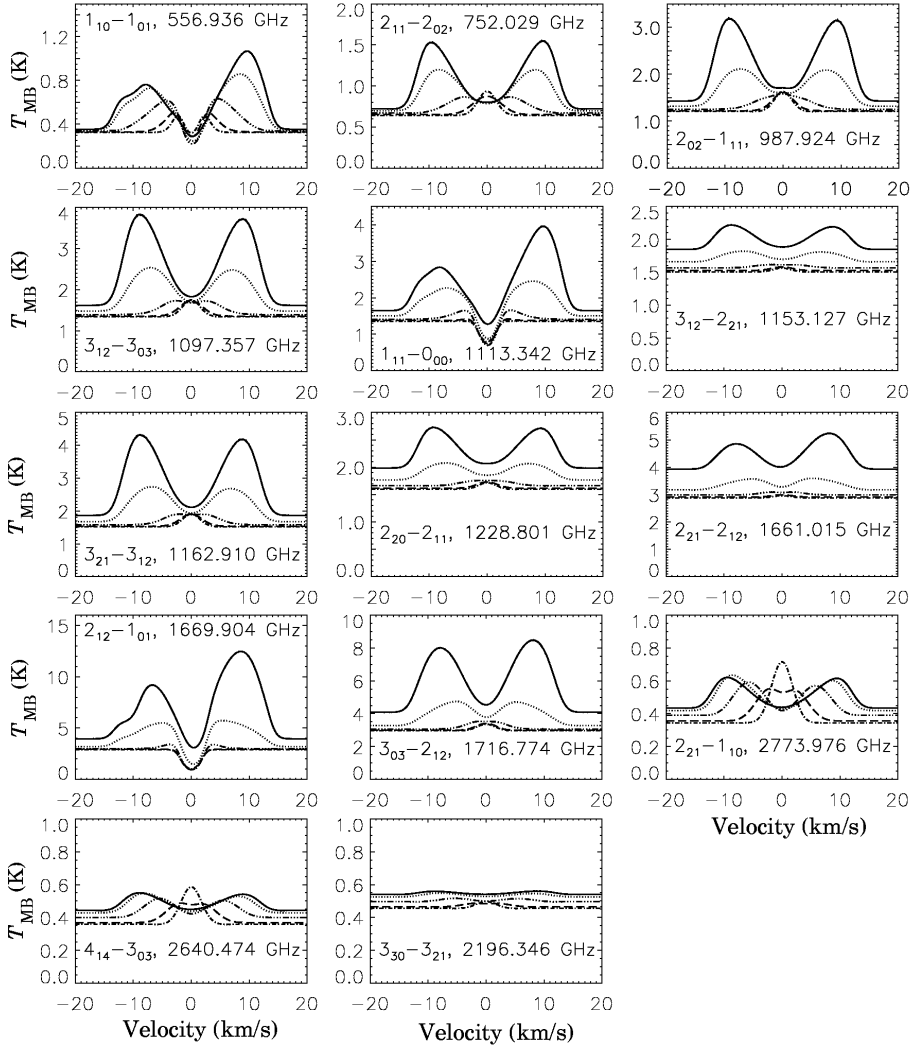


Figure 5.5: Line profiles of H_2O for a model with $\alpha = 1$, a cavity with half-opening angle $\theta = 15^\circ$ and an outflow velocity field with $V_{\text{out}} = 12 \text{ km s}^{-1}$. The abundance outside the cavity has a constant value of 10^{-9} , whereas inside the cavity the abundance is 10^{-7} . The different line profiles are the result of the different inclination angles. Inclination of 0° means face-on, as the line of sight is directed along the cavity; inclination of 90° is edge-on. Inclination = 0° (solid), 20° (dotted), 45° (dash-dotted), 70° (dashed), 90° (solid).

$n(\text{H}_2) = 5 \times 10^3 \text{ cm}^{-3}$, in agreement with the overall density found through hydrodynamical modeling of outflows (Shang et al. 2006), and a coupled gas and dust temperature of 300 K. This high gas and dust temperature is reasonable as in outflows the temperatures reach a few 100 K (Giannini et al. 2001). Due to these high temperatures, enhanced neutral-neutral rates lead to elevated H_2O abundances. For this reason, a constant water abundance of 10^{-7} is chosen throughout the outflow. The outcome of models with a water abundance of 10^{-6} in the outflow is discussed in Section 5.5. The outflow velocity field is modeled as a function of radius, i.e., $V \propto R$, reaching a maximum velocity of $\pm 12 \text{ km s}^{-1}$ at the outer radius. Resulting line profiles are plotted in Fig. 5.5, integrated line strengths are listed in Table 5.3.

One notices the difference in width of the line profiles for models with an inclination angle of 0° compared to 90° . This is the result of the fraction of the outflow velocity field lying in the line of sight, i.e., face-on models tend to have a broader line profile as opposed to edge-on models. It is seen from Table 5.3 that a decrease in inclination angle leads to an increase in the line strengths as the cavity is more uncovered. The outflow is best seen in the ground-state transitions for models with a small inclination angle, i.e., close to face-on. This is expected since the low density cone favours the emission

Table 5.3: Integrated intensities, $\int T_{\text{MB}} \Delta v$ [K km s^{-1}], for a model with an outflow cavity with half-opening angle $\theta = 15^\circ$.

Transition	Frequency [GHz]	Inclination angle				
		0°	20°	45°	70°	90°
ortho- H_2O transitions						
$1_{10}-1_{01}$	557.936	8.7	6.8	3.1	1.1	0.7
$3_{12}-3_{03}$	1097.35	32.9	17.4	3.7	1.8	1.6
$3_{12}-2_{21}$	1153.12	5.5	2.7	0.5	0.3	0.2
$3_{21}-3_{12}$	1162.91	36.6	17.8	3.4	1.7	1.5
$2_{21}-2_{12}$	1661.01	17.5	7.0	0.8	0.4	0.3
$2_{12}-1_{01}$	1669.90	105.8	33.9	-5.0	-7.2	-6.8
$3_{03}-2_{12}$	1716.17	67.2	25.3	3.6	1.8	1.5
$3_{30}-3_{21}$	2196.34	0.3	0.3	0.3	0.2	0.2
$4_{14}-3_{03}$	2640.47	1.6	1.6	1.4	1.1	1.0
$2_{21}-1_{10}$	2773.97	2.7	2.8	2.5	1.9	1.7
para- H_2O transitions						
$2_{11}-2_{02}$	752.029	12.1	8.1	2.7	1.3	1.2
$2_{02}-1_{11}$	987.924	26.6	13.7	3.1	1.7	1.6
$1_{11}-0_{00}$	1113.34	25.6	12.2	-0.3	-1.8	-1.7
$2_{20}-2_{11}$	1228.80	11.2	5.3	1.0	0.5	0.5

from low-lying transitions, in that the critical densities of the water lines are of the order of 10^8 cm^{-3} . From Fig. 5.5 we see that the peak in the blue wing is weaker than the peak in the red wing of the line profile for the $1_{10}\text{-}1_{01}$ ground-state transition, in contrast with the line profiles of the constant abundance and jump profile models. Note that for an inclination angle of 0° , the optical depth through the center, i.e., along the outflow, is larger than 1 for the ground-state transition. It is found that models with an inclination angle of 90° produce line profiles similar to the profiles in the case of a model with a constant water abundance of 10^{-9} , except for the transitions seen with ISO. However, line strengths differ by a factor of a few.

Clearly, HIFI's sensitivity can put constraints on the inclination angle of the outflow, in that for small inclination angles, the blue wing is found to be less strong than the red wing.

5.4.4 Circumstellar disk

Even though the formation of high-mass stars is not completely understood, more and more observational evidence is found that the creation originates in a similar manner as its low-mass counterpart (Cesaroni et al. 1999). Therefore, we examine the influence of an embedded disk in the center of the AFGL 2591 region on the line profiles and strengths. We consider the scenario of Section 5.4.1, for which we adopt a constant water abundance of 10^{-9} throughout the region.

The size of a disk can be quantified as $r_d = 350 \text{ AU} (\gamma/0.007) (M_*/20 M_\odot)^{1/2}$, where γ is the ratio of rotational to gravitational energy. For high-mass star formation, $\gamma \sim 0.007$ (Pirogov et al. 2003). When adopting a mass of the central star toward AFGL 2591 of $16 M_\odot$, as inferred by van der Tak & Menten (2005), a disk with radius $\sim 300 \text{ AU}$ is found. We adopt an overall disk density of 10^9 cm^{-3} , and a water abundance of 10^{-6} . The temperature of the gas and dust is assumed to be 100 K . Note that the characterization of the density and temperature structure of the disk is far from accurate. However, the intent of this section is to investigate the influence of a 'simple' disk model, not to model a realistic disk. The excitation of rotational, as well as ro-vibrational, water transitions in a circumstellar disk when the chemistry and temperature structure is calculated self-consistently is discussed in chapter 6 (Meijerink et al. 2007).

It is found that the line profiles and strengths are insignificantly influenced under all possible inclination angles (less than 1°). Since HIFI will be able to spectrally resolve disk structures, the emission from our adopted disk model with a diameter up to a few hundred AU in size is overwhelmed by that of the envelope within the HIFI beams. However, observations show that more massive, elongated disks with a diameter up to few thousands of AU are likely to exist (Beltrán et al. 2005). Due to high densities

and large columns of dust in circumstellar disks, the stellar radiation field is attenuated. As a result, the temperature of the gas and dust at radial distances of a more than a few 10 AU drops below 100 K (see Glassgold et al. 2007). Consequently, water will be frozen out on dust grains, reducing significantly the gas phase water abundance. When adopting the same disk density as in previous model, a water abundance value of 10^{-8} , and a disk radius of 1500 AU, the line profiles differ only a few % from the models without a disk.

5.4.5 Predictions for PACS

Besides HIFI detections, observations by PACS will enlarge significantly our knowledge on high-mass star-forming regions. PACS will carry out spectral scans from 57 to 210 μm , thereby providing information on the dust continuum emission at wavelengths currently poorly characterized. However, the PACS data will be spectrally unresolved ($\lambda/\Delta\lambda \sim 1500$), in contrast to HIFI observations.

In this section, we compute the spectral output for PACS for two different scenarios. The first scenario (Model I) is for a model with a constant water abundance of 10^{-8} , the second scenario (Model II) is for the model as described in Section 5.4.3, in which the outflow cavity is seen under an inclination angle of 0° . Resulting spectra are plotted in Fig. 5.6, integrated line strengths are listed in Table 5.4.

It is seen from Fig. 5.6 that the two models lead to rather different output spectra. Strong lines are $2_{20}-1_{11}$, $2_{21}-1_{10}$, $4_{14}-3_{03}$, $3_{03}-2_{12}$, and $2_{12}-1_{01}$. Depending on the model, the $3_{21}-2_{12}$, $2_{20}-1_{11}$, $2_{21}-1_{10}$, and $2_{12}-1_{01}$ transitions are found either in emission or in absorption. One notices the change in the behaviour of the dust continuum emission between the two models. Hence, PACS spectra will tell us about the geometry of these regions, both through the continuum and through the lines. Note that the signal-to-noise (S/N) ratios for the different lines vary between 30 and 3000 for Model I and between 25 and 300 for Model II.

It is now interesting to compare the results from Kaufman & Neufeld (1996) in which far-infrared water emission from shock waves is presented. They find that the strongest transitions are the $2_{12}-1_{10}$, $3_{03}-2_{12}$, $2_{21}-1_{10}$, $4_{14}-3_{03}$, $3_{13}-2_{02}$, $1_{10}-1_{01}$, and $3_{22}-2_{11}$ lines. We find that these transitions are also the strongest in the case of models with an outflow cavity.

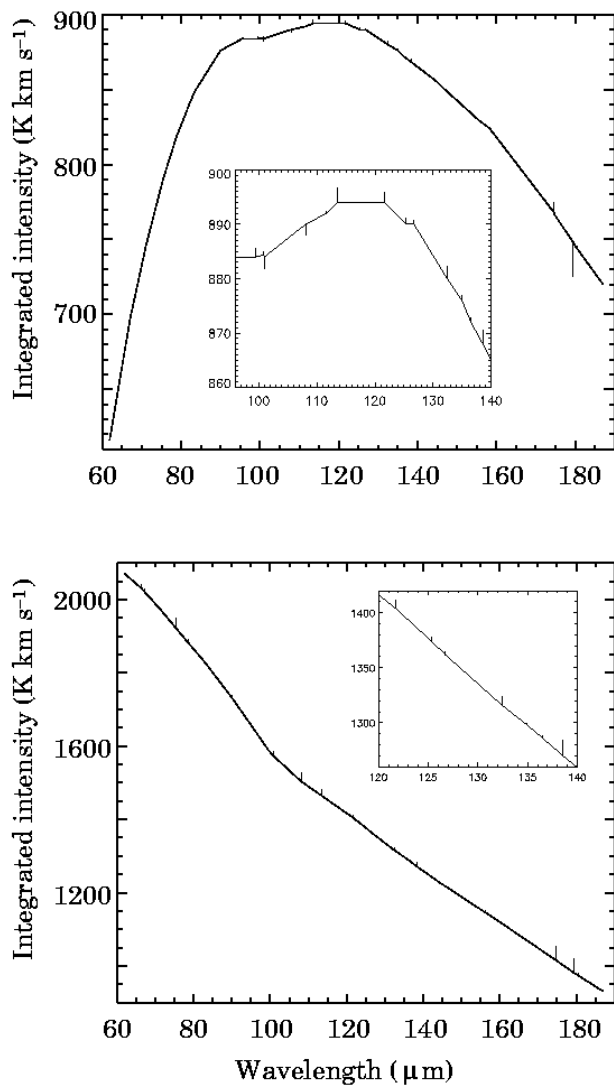


Figure 5.6: Predictions for the spectrum to be observed with PACS for two different scenarios. Integrated intensities are plotted on top of the continuum background. The top panel denotes the spectrum for a model with a constant water abundance of 10^{-8} . Bottom panel plots the spectrum when a cavity with an outflow velocity field, with $V_{\text{out}} = 12 \text{ km s}^{-1}$, is seen under an inclination angle of 0° . The Y-axis is in K km s^{-1} . To convert to W cm^{-2} , multiply by $v^3 T_{\text{MB}} \Delta v \times 10^{-57}$.

Table 5.4: Integrated intensities, $\int T_{\text{MB}} \Delta \nu$ [K km s^{-1}], for two different models, when convolved with the PACS beam. In Model I a constant water abundance of 10^{-8} is assumed. Model II describes AFGL 2591 as a region with a cavity outflow, with $V_{\text{out}} = 12 \text{ km s}^{-1}$, seen face-on, i.e., inclination angle = 0° . To convert to W cm^{-2} , multiply by $\nu^3 T_{\text{MB}} \Delta \nu \times 10^{-57}$.

Transition	Wavelength [μm]	Model I	Model II
ortho- H_2O transitions			
3 ₃₀ -2 ₂₁	66.48	0.3	13.2
3 ₂₁ -2 ₁₂	75.43	-1.7	27.0
4 ₂₃ -3 ₁₂	78.79	0.6	11.7
5 ₀₅ -4 ₁₄	99.56	1.5	4.3
2 ₂₁ -1 ₁₀	108.1	-1.9	25.4
4 ₁₄ -3 ₀₃	113.5	2.7	17.6
4 ₃₂ -4 ₂₃	121.7	1.8	7.6
4 ₂₃ -4 ₁₄	132.4	2.2	7.9
3 ₀₃ -2 ₁₂	174.6	6.9	36.5
2 ₁₂ -1 ₀₁	179.5	-23.1	39.2
para- H_2O transitions			
3 ₃₁ -2 ₂₀	67.13	0.2	8.3
3 ₂₂ -2 ₁₁	89.99	0.4	7.6
2 ₂₀ -1 ₁₁	100.9	-2.2	14.8
3 ₁₃ -2 ₀₂	138.5	2.5	14.4

5.5 Discussion

5.5.1 Comparison to SWAS and ISO observations

Observations by SWAS of the ground-state transition of o- H_2O at 557 GHz reveal an asymmetric line profile (Boonman et al. 2003). However, radiative transfer modeling of line profiles in the massive young stellar object AFGL 2591 show that constant and jump water profile distributions do not fit this SWAS observation of the ground-state transition of o- H_2O , in that a symmetric line profile is found. In contrast to these models, an asymmetric line profile is derived for the ground-state transition for models with an outflow cavity and an inclination angle between 0° and 30° . This agrees with the observations by van der Tak et al. (2006), who conclude that the outflow is seen under an inclination angle of 26° – 38° . The integrated intensity of ± 7 – 9 K km s^{-1} reduces, after convolution with a beam size comparable to the size of the SWAS beam, i.e., $3.3' \times 4.5'$, to a value of $\pm 0.5 \text{ K km s}^{-1}$. This

value is in agreement with the observed $0.37 \pm 0.04 \text{ K km s}^{-1}$ strength. Note that this observed value is a lower limit, and that the actual value may be increased by a factor of ~ 2 , due to the presence of foreground clouds as discussed in Boonman et al. (2003).

Table 5.5 lists the observed integrated line intensities for some selected ISO transitions, published in Boonman et al. (2003). The observed line fluxes [$\text{W cm}^{-2} \mu\text{m}^{-1}$] have been converted into integrated line intensities [K km s^{-1}], by adopting the appropriate ISO beam size, to simplify comparison with the results presented in this chapter. For the 3_{30} - 3_{21} and 4_{14} - 3_{03} transitions, the predicted integrated intensities are below the observed strengths for all the models by a factor of ~ 4 , except for the models with an outflow that overproduce the 4_{14} - 3_{03} strength by a factor of ~ 2 . The 2_{21} - 1_{10} line is observed in absorption, whereas our models favour to produce this transition in emission. The model with a constant water abundance of 10^{-7} is favoured to fit this transition. However, due to the lack of more o- and p- H_2O observations and the poor spectral resolution of the ISO satellite, the ISO-LWS data are not decisive on which model is favoured. The Herschel mission will bring additional (line profile) information on water transitions with somewhat lower energies, which is needed to decide between the different models.

Table 5.5: Observed ISO and SWAS line intensities [K km s^{-1}] for AFGL 2591, see Boonman et al. 2003

Transition and Wavelength [GHz]			
2773.97	2640.47	2196.34	557.936
2_{21} - 1_{10}	4_{14} - 3_{03}	3_{30} - 3_{21}	1_{10} - 1_{01}
< -0.5	< 0.7	< 1.2	0.37 ± 0.04

5.5.2 Parameter dependency

Steepness of density gradient

We examine the influence of the adopted density and temperature distribution, i.e., adopting an $\alpha = 2$ profile, on the resulting line profiles and strengths. In analogy to the models described in Section 5.4.1, we assume the water abundance, ranging from 10^{-10} to 10^{-7} , to be uniformly distributed. Resulting line profiles are plotted in Fig. 5.7, integrated brightness temperatures are listed in Table 5.5.2.

In general, the line shapes are similar to the profiles presented in Section 5.4.1. However, the 2_{21} - 2_{12} , 3_{03} - 2_{12} , and the ISO transitions are found in absorption, in contrast to the $\alpha = 1$ models. This is a consequence of a

steeper density slope, providing more material within the Herschel beam. Note that the optical depth in the lines increases by a factor of $\sim 5-10$ for the high-lying transitions. The line strengths of the ground-state lines, i.e., $1_{10}-1_{01}$ and $1_{11}-0_{00}$, are similar to the strengths for the $\alpha=1$ models. This is not surprising since in the outer regions the temperature distribution follows the distribution of the $\alpha=1$ model.

In addition, the effect of an $\alpha=2$ temperature and density profile on the line profiles and strengths is tested for models with an outflow cavity, described in Section 5.4.3. We find that the line shapes are similar to the profiles for the models with an $\alpha=1$ density gradient. However, the line strengths drop by a factor of ~ 2 to 3 for the different transitions.

Opening angle

In this section we briefly discuss the influence of the adopted outflow opening angle. We adopt density, temperature and abundance distributions from Section 5.4.3. Outflows with a cavity half-opening angle of 7.5° and 22.5° are computed. It is found that the adopted opening has minor influ-

Table 5.6: Integrated intensities, $\int T_{\text{MB}} \Delta v$ [K km s^{-1}], for models with a constant H_2O abundance. The adopted water abundance is denoted on top of each column. The density and temperature distribution is described by a $\alpha=2$ profile coefficient.

Transition	Frequency [GHz]	10^{-10}	10^{-9}	10^{-8}	10^{-7}
ortho- H_2O transitions					
$1_{10}-1_{01}$	557.936	-6.0(-2)	-0.6	-1.6	-2.1
$3_{12}-3_{03}$	1097.35	5.6(-2)	0.6	3.5	9.9
$3_{12}-2_{21}$	1153.12	1.8(-2)	0.2	2.4	15.4
$3_{21}-3_{12}$	1162.91	0.1	1.0	7.0	23.7
$2_{21}-2_{12}$	1661.01	-0.2	-1.6	-8.2	-18.3
$2_{12}-1_{01}$	1669.90	-2.3	-14.5	-31.9	-42.4
$3_{03}-2_{12}$	1716.77	-0.1	-1.1	-6.4	-16.8
$3_{30}-3_{21}$	2196.34	-1.1(-3)	-1.0(-2)	-6.1(-2)	-0.1
$4_{14}-3_{03}$	2640.47	-3.4	-3.1(-2)	-0.2	-0.4
$2_{21}-1_{10}$	2773.97	-7.5(-3)	-6.8(-2)	-0.4	-0.9
para- H_2O transitions					
$2_{11}-2_{02}$	752.029	3.1(-2)	0.3	2.4	8.6
$2_{02}-1_{11}$	987.924	7.2(-2)	0.7	5.2	8.7
$1_{11}-0_{00}$	1113.34	-0.4	-3.2	-10.1	-14.5
$2_{20}-2_{11}$	1228.80	3.0(-2)	0.3	2.2	6.3

NOTE: $a(b)$ means $a \times 10^b$

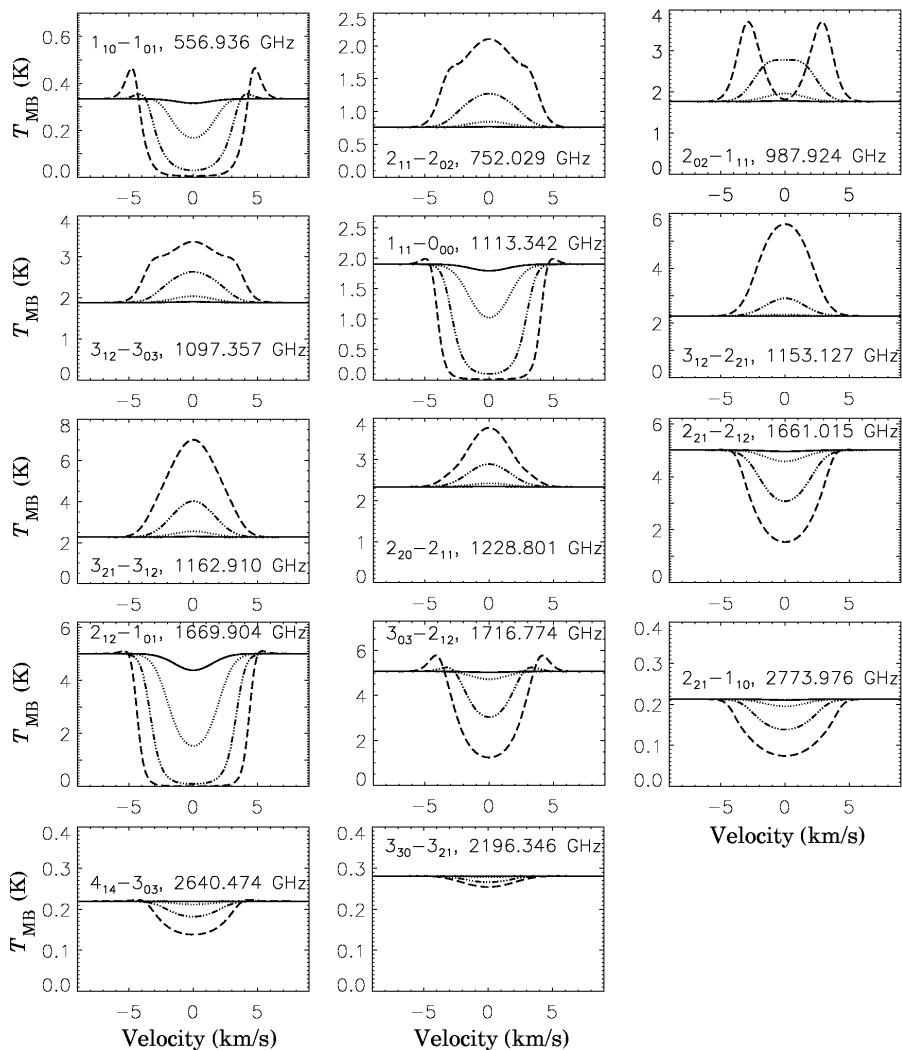


Figure 5.7: Line profiles of ortho- and para-H₂O transitions for models with $\alpha = 2$. The water abundance, i.e., o-H₂O + p-H₂O, is uniformly distributed across the region and ranges from 10^{-10} to 10^{-7} , with 10^{-10} (solid), 10^{-9} (dotted), 10^{-8} (dash-dotted), 10^{-7} (dashed).

ence on the line shapes. However, the line strengths drop (increase) by a factor of $\sim 2-3$ (2) for models with an half-opening angle of 7.5° (22.5°).

Outflow water abundance

In this section we briefly discuss the influence of an increased outflow water abundance by a factor of 10, i.e., 10^{-6} . We adopt density, temperature and abundance distributions from Section 5.4.3. It is found that the adopted outflow water abundance has minor influence on the line shapes. However, the line strengths increase by a factor of ~ 2 to 10. Note that the lowest gain in the line strengths is found for the ground-state transitions, i.e., $1_{10}-1_{01}$, $2_{12}-1_{01}$, and $1_{11}-0_{00}$, and that the line strengths of the higher-lying transitions, e.g., $3_{12}-3_{03}$, $3_{12}-2_{21}$, $3_{21}-3_{12}$, $2_{21}-1_{01}$ and $2_{20}-2_{11}$, are increased by a factor up to ~ 10 . This difference in enhancement factor can be explained due to the high gas and dust temperature in the outflow, favouring the excitation of the high-lying transitions.

5.6 Summary and conclusions

We have constructed models to examine the excitation of water in the high-mass star-forming region AFGL 2591. Depending on the adopted density, temperature and abundance structure, a completely different set of line profiles and strengths is found. Hence, the line profiles are very sensitive to the adopted physical and chemical structure. We have found that (i) the ground-state transitions $1_{10}-1_{01}$, $2_{12}-1_{01}$ and $1_{11}-0_{00}$, with relatively low upper energy levels ($\lesssim 110$ K), become highly optically thick in the outer regions. The line profiles for these transitions, are mainly dominated by the emission from the outer regions, and are therefore not useful to put constraints on the water abundance in the inner regions. However, (ii) the emission from lines with higher upper energy levels is dominated by the emission originating in the inner regions, and are therefore useful to probe the water abundance in the warm inner regions. (iii) For models with an outflow cavity, the outflow feature (blue peak less strong than the red peak) is best seen in the ground-state transitions of o- and p- H_2O . (iv) The influence of a moderate disk (few 100 AU in size) in the centre of the AFGL 2591 region does not change the water line profiles and strengths within the Herschel beam.

The Herschel mission will thus greatly help to understand the structure of high-mass protostellar objects, and consequently the formation process of high-mass stars.

Acknowledgements

We thank Xander Tielens, Marco Spaans, Tim van Kempen and Ewine van Dishoeck for helpful discussions and suggestions which have improved the chapter.

References

- Arce, H. G., Shepherd, D., Gueth, F., et al. 2007, in *protostars and Planets V*, ed. B. Reipurth, D. Jewitt, & K. Keil, 245-260
- Bachiller, R., & Perez Gutierrez, M. 1997, *ApJ*, 487, L93
- Bally, J., & Lada, C. J. 1983, *ApJ*, 265, 824
- Beech, M., & Mitalas, R. 1994, *ApJS*, 95, 517
- Beltrán, M. T., Cesaroni, R., Neri, R., et al. 2004, *ApJ*, 601, L187
- Beuther, H., Schilke, P., Sridharan, T. K., et al. 2002, *A&A*, 383, 892
- Bonnell, I. A., & Bate, M. R. 2002, *MNRAS*, 336, 659
- Bonnell, I. A., Bate, M. R., & Zinnecker, H. 1998, *MNRAS*, 298, 93
- Boonman, A. M. S., Doty, S. D., van Dishoeck, E. F., et al. 2003, *A&A*, 406, 937
- Boonman, A. M. S., & van Dishoeck, E. F. 2003, *A&A*, 403, 1003
- Campbell, B. 1984, *ApJ*, 287, 334
- Cesaroni, R. 2005, *Ap&SS*, 295, 5
- Cesaroni, R., Felli, M., Jennes, T., et al. 1999, *A&A*, 345, 949
- de Graauw, T., Haser, L. N., Beintema, D. A., et al. 1996, *A&A*, 315, L49
- Doty, S. D., & Neufeld, D. A. 1997, *ApJ*, 489, 122
- Doty, S. D., van Dishoeck, E. F., & Tan, J. C. 2006, *A&A*, 454, L5
- Egan, M. P., Leung, C. M., & Spagna, G. F. 1998, *Computer Physics Communications*, 48, 271
- Elitzur, M. 1992, *Astronomical masers (Astronomical masers Kluwer Academic Publishers (Astrophysics and Space Science Library. Vol. 170), 365 p.)*
- Giannini, T., Nisini, B., & Lorenzetti, D. 2001, *ApJ*, 555, 40
- Glassgold, A. E., Najita, J. R., & Igea, J. 2007, *ApJ*, 656, 515
- Hasegawa, T. I. & Mitchell, G. F. 1995, *ApJ*, 451, 225
- Helmich, F. P., van Dishoeck, E. F., Black, J. H., et al. 1996, *A&A*, 315, L173
- Hogerheijde, M. R., & van der Tak, F. F. S. 2000, *A&A*, 362, 697
- Kaufman, M. J., & Neufeld, D. A. 1996, *ApJ*, 456, 611
- Kessler, M. F., Steinz, J. A., Anderegg, M. E., et al. 1996, *A&A*, 315, L27
- Klessen, R. 2001, in *ASP Conf. Ser. 243: From Darkness to Light: Origin and Evolution of Young Stellar Clusters*, ed. T. Montmerle & P. André, 139
- Klessen, R. S., Spaans, M., & Jappsen, A.-K. 2007, *MNRAS*, 374, L29
- Krumholz, M. R., McKee, C. F., & Klein, R. I. 2005, *ApJ*, 618, L33
- Lada, C. J., Thronson, Jr., H. A., Smith, H. A., Schwartz, P. R., & Glaccum, W. 1984, *ApJ*, 286, 302
- McKee, C. F., & Tan, J. C. 2002, *Nature*, 416, 59
- McKee, C. F., & Tan, J. C. 2003, *ApJ*, 585, 850
- Meijerink, R., Poelman, D. R., Spaans, M., Glassgold, E. A., & Tielens, A. G. G. M. 2007, in preparation
- Mitchell, G. F., Hasegawa, T. I., & Schella, J. 1992, *ApJ*, 386, 604
- Mumma, M. J., Weaver, H. A., & Larson, H. P. 1987, *A&A*, 187, 419
- Ossenkopf, V., & Henning, T. 1994, *A&A*, 291, 943
- Palla, F., & Stahler, S. W. 1993, *ApJ*, 418, 414
- Pilbratt, G. L. 2005, in *The Dusty and Molecular Universe: A Prelude to Herschel and ALMA*, ed. A. Wilson, 3-10
- Pirogov, L., Zinchenko, I., Caselli, P., Johansson, L. E. B., & Myers, P. C. 2003, *A&A*, 405, 639
- Poelman, D. R., & Spaans, M. 2005, *A&A*, 440, 559

- Poelman, D. R., & Spaans, M. 2006, *A&A*, 453, 615
- Sandell, G., Wright, M., & Forster, J. R. 2003, *ApJ*, 590, L45
- Schneider, N., Bontemps, S., Simon, R., et al. 2006, *A&A*, 458, 855
- Schöier, F. L., van der Tak, F. F. S., van Dishoeck, E. F., & Black, J. H. 2005, *A&A*, 432, 369
- Shang, H., Allen, A., Li, Z.-Y., et al. 2006, *ApJ*, 649, 845
- Shepherd, D. 2003, in *ASP Conf. Ser. 287: Galactic Star Formation Across the Stellar Mass Spectrum*, ed. J. M. De Buizer & N. S. van der Blik, 333-344
- Shepherd, D. 2005, in *IAU Symposium*, ed. R. Cesaroni, M. Felli, E. Churchwell, & M. Walm-
sley, 237-246
- Snell, R. L., Howe, J. E., Ashby, M. L. N., et al. 2000, *ApJ*, 539, L101
- Steinacker, J., Chini, R., Nielbock, M., et al. 2006, *A&A*, 456, 1013
- Trinidad, M. A., Curiel, S., Cantó, J., et al. 2003, *ApJ*, 589, 386
- van der Tak, F., Neufeld, D., Yates, J., et al. 2005, in *The Dusty and Molecular Universe: A
Prelude to Herschel and ALMA*, ed. A. Wilson, 431-432
- van der Tak, F. F. S., & Menten, K. M. 2005, *A&A*, 437, 947
- van der Tak, F. F. S., van Dishoeck, E. F., Evans, II, N. J., Bakker, E. J., & Blake, G. A. 1999,
ApJ, 522, 991
- van der Tak, F. F. S., van Dishoeck, E. F., Evans, II, N. J., & Blake, G. A. 2000, *ApJ*, 537, 283
- van der Tak, F. F. S., Walmsley, C. M., Herpin, F., & Ceccarelli, C. 2006, *A&A*, 447, 1011
- van Dishoeck, E. F. & Helmich, F. P. 1996, *A&A*, 315, L177
- van Zadelhoff, G.-J., Dullemond, C. P., van der Tak, F. F. S., et al. 2002, *A&A*, 395, 373
- Wolfire, M. G., & Cassinelli, J. P. 1987, *ApJ*, 319, 850
- Yorke, H. W., & Krügel, E. 1977, *A&A*, 54, 183
- Yorke, H. W., & Sonnhalter, C. 2002, *ApJ*, 569, 846
- Zhang, Q., Hunter, T. R., Sridharan, T. K., & Ho, P. T. P. 2002, *ApJ*, 566, 982